Introduction to the XMM-Newton Science Analysis System

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Outline

- What is the SAS?
- SAS installation and setup
- SAS data reduction scheme
- Getting started I: the Observation Data File (ODF) + odfbrowser
- Getting started II: the Calibration Current File (CCF) and the Calibration Index File (CIF)
- Getting started III: the common first steps
- Running SAS: GUI or command line
- PPS or “SAS has already reduced these data”
- XSA can reprocess the data (RISA)
- Retrieving XMM-Newton data
- SAS >> web services - shaping the future

What is the SAS?

- The XMM-Newton Scientific Analysis System is a suite of programs (“tasks”) for dealing with data from all XMM-Newton Instruments
- It is written basically in C++ and Fortran 90/95. Perl and shell scripts constitute “metatasks”. It makes use of public libraries / programs like cfitsio, Xmgrace, ds9
- It has been developed by ~ 30 programmers, working in 6 different countries
- A subset of the SAS is used as the core of the official pipeline (PPS) for reducing the data to calibrated event lists, images, spectra, source lists (and much more)

Actual SAS version: SAS 17 – released on June 22, 2018

Installation and setup

- Binary distributions of SAS are available **only** for 64bit:
  - Linux RedHat 6.8, Ubuntu 16.04
  - Mac OS X - 64bit
    + 10.12.6 - Sierra, can be used with 10.11 - El Capitan
    or 10.13 High Sierra (not fully tested yet with Mojave, but seems to work)
  - *Virtual Machine running Ubuntu 18.04 64bit (Windows, Linux and MacOSX)...*

  >> "official" supported platforms which can be used by other OSs

- Objectives:
  - make it easy to install: untar and go
  - provide all libraries required, also external ones (like cfitsio) - (however, need to be installed: ds9 / FTOOLS / GRACE / Perl)

Installation and setup II

All SAS installations are binary (no support for building from source code)

tar zxf sas_17.0.0-[OS].tgz

./install.sh (will check everything is in place, download and install a Miniconda if necessary and finally install SAS)

>> xmmsas_20180620_1731 directory with all contents

Everytime you want to run SAS:

>. ./setsas.sh (bash) or > source ./setsas.csh (csh) in that directory

or

> . <top-dir>/xmmsas_20180620_1731/setsas.sh
> source <top-dir>/xmmsas_20180620_1731/setsas.csh

Basics of X-ray data

X-ray detectors are photon-counting → two main consequences:

- X-ray astronomy is an intrinsic Poissonian science
  - Scientific products can have a few or even zero events in large ranges of their parameter spaces

- The "king" in the X-ray realm is the event, characterised by:
  - **position (X-Y) on the detector**
  - **"pulse height", which is related to the X-ray energy (E) of the incoming photon in a complex and generally non-linear way**
  - **arrival time (t) at the spacecraft**
  - event "shape" (used to separate X-ray events from particles' signatures)
  - other secondary attributes (you don't generally have to worry about)

Event lists

When? **Where?** **Who?** **What?**

	TIME	X	Y	PHA	PI	PATTERN	CCDNR
	D	J	J	I	I	B	B
	s	0.05 ARCSECONDS	0.05 ARCSECONDS	CHAN	CHAN		
1	9.506202266412E+07	23743	21330	423	1447	2	1
2	9.506202266412E+07	28728	21990	25	98	0	1
3	9.506202527717E+07	28176	31623	25	97	0	1
4	9.506202527717E+07	29829	30841	327	1131	0	1
5	9.506202527717E+07	23686	19319	541	1854	0	1
6	9.506203046611E+07	25510	32711	1810	6171	0	1
7	9.506203566620E+07	29814	28823	102	360	0	1
8	9.506203826626E+07	26635	30601	2062	7028	0	1
9	9.506204346625E+07	26429	20314	443	1519	4	1
10	9.506204606629E+07	20691	28728	1608	5471	3	1
11	9.506204606629E+07	27989	29777	202	700	0	1
12	9.506204606629E+07	21937	25667	117	402	2	1
13	9.506204866632E+07	28132	32491	462	1589	0	1
14	9.506204866632E+07	27204	29741	904	3095	0	1
15	9.506205126638E+07	22124	20257	290	994	0	1
16	9.506205906643E+07	23193	18795	1398	4771	0	1
17	9.506206166646E+07	23224	19326	276	950	0	1
18	9.506206946653E+07	27755	28979	183	637	0	1
19	9.506207206939E+07	22533	29563	33	118	0	1

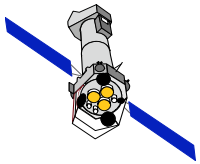
Science as projection

The X-ray scientific products can be seen as ***projections*** onto the sub-spaces defined by the event physical quantities

- By collapsing time and space, one gets an energy distribution function (***spectrum***) in units of ***counts per energy bin***
- By collapsing time and energy, one gets a 2-D ***image*** in units of ***counts per pixel***
- By collapsing space and energy, one gets an intensity ***time series*** in units of ***counts per time bin***

These scientific products are expressed in units that are *in*directly related to the intrinsic properties of celestial sources

Data reduction = calibration



Time calibration

Optics calibration

Image analysis:
- Exposure maps
- *Point Spread Function*

Timing analysis:
- Exposure maps
- *Point Spread Function*

ODF

ODF reduction

Scientific product extraction

Scientific product analysis

Science

Attitude calibration

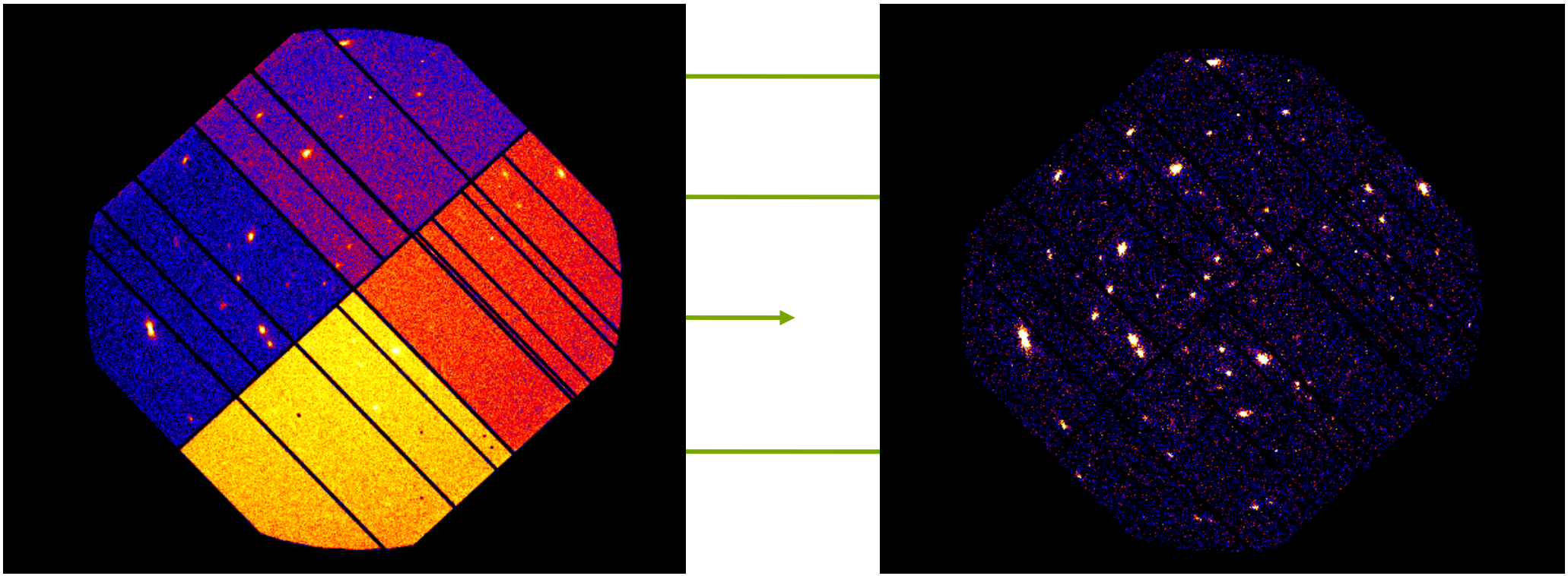
Detector calibration

Energy spectra analysis:
- Effective area
- Redistribution matrix
- *Line Spread Function*

Reduction scope

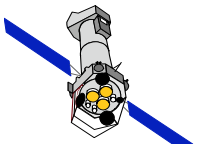
SAS does two things (to XMM data), that no other tool does:

- applies calibrations to raw data



- optimally screen / filter your data

SAS Grand-Scheme

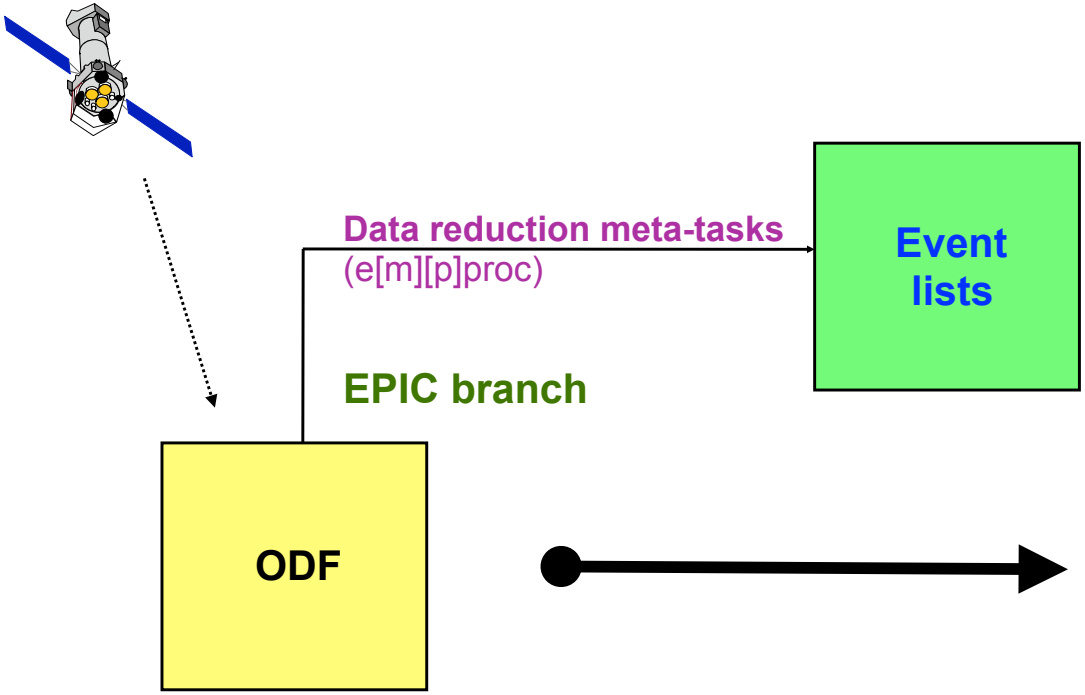


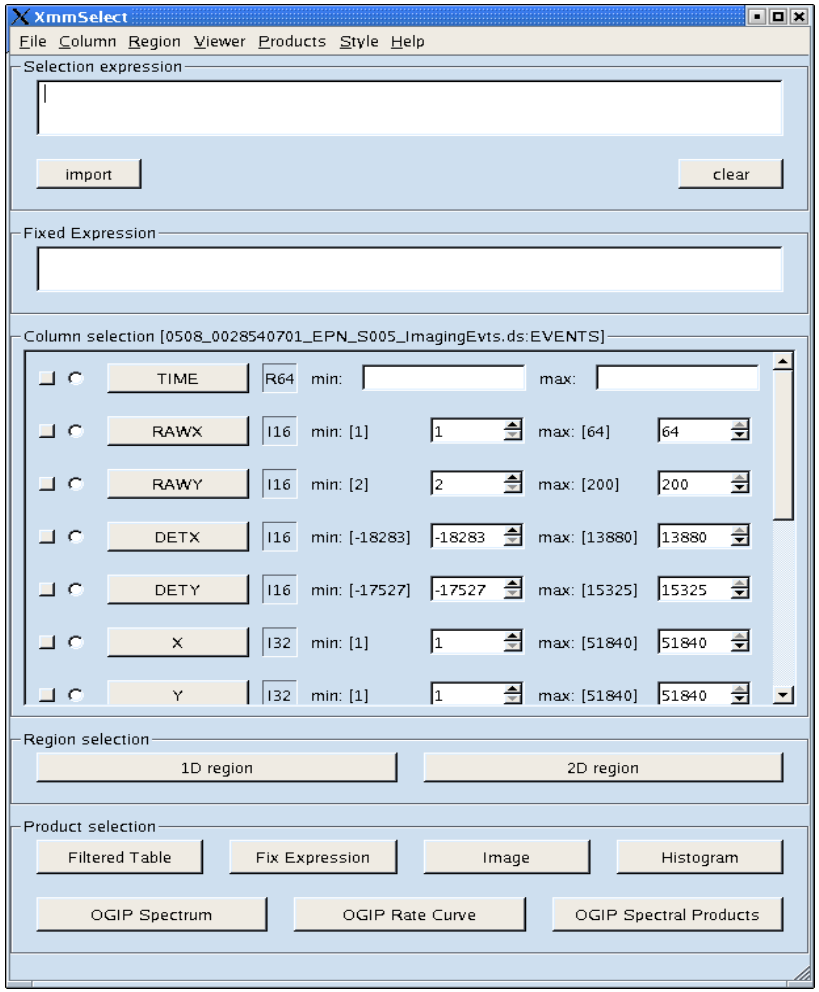
ODF contents

ODF

- CCD-based event lists, containing uncalibrated quantities
- Auxiliary and Housekeeping files, pn/RGS diagnostic images
- Spacecraft housekeeping — FITS
- Spacecraft attitude showing the satellite star tracker pointing
- Time correlation file (onboard time and frame counter versus UTC)
- ODF summary file — ASCII

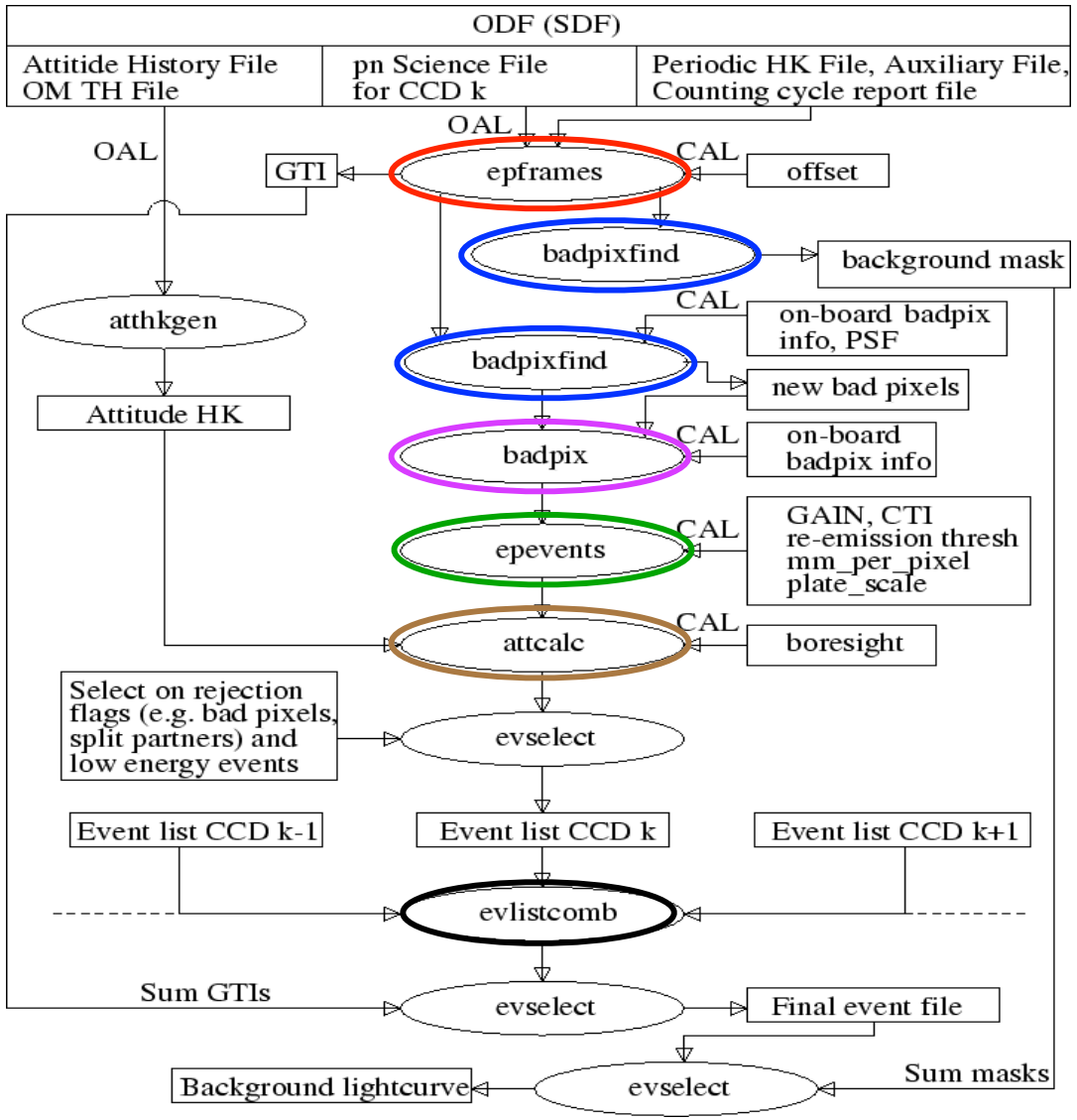
epicproc = e[m][p]proc





- metatasks to process MOS or pn data
- generate calibrated, filtered event lists
- leave user in control of GTI and filter expressions

epproc reduction scheme

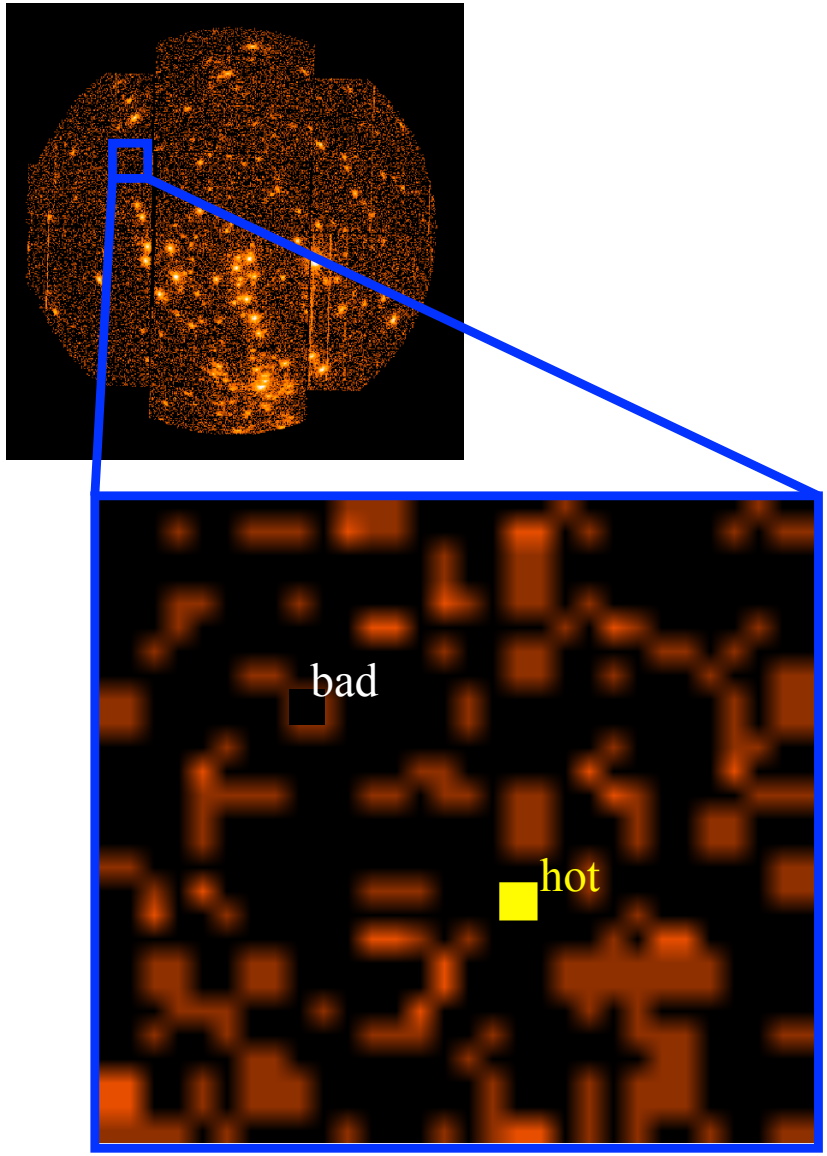


- **epframes** to process a CCD, exposure and datamode specific ODF file, creating the output raw event list and GTI data set
- **badpixfind** to find new bad pixels
- **badpix** to process the raw event list, adding the BADPIX extension
- **epevents** to process the event list file, flagging trailing events, performing split events pattern recognition, CTI and gain correction to create the calibrated event list
- **attcalc** to calculate the X and Y sky coordinates.
- **evlistcomb**, the CCD specific data sets are merged into a single event list.

Event lists

- In the EPIC pn imaging mode, the EVENTS table of the calibrated event list files contain 14 columns i.e :
  - TIME --> when did my photon arrive
  - RAWX RAWY --> where on the CCD
  - DETX DETY -->where on the detector
  - X Y --> where from the sky
  - PHA PI --> which energy did my photon have
  - FLAG --> did it hit the detector at a critical place
  - PATTERN --> was it a single/double.....
  - CCDNR --> on which CCD did it hit the detector

Bad pixels



- dead pixel: no events are detected
- hot pixel: pixel "produces" ghost events very often
- by default epicproc will try to detect bad pixels for any imaging exposure.
- the new bad pixels are then used in the data reduction together with any other known (via the calibration files) bad pixels

Data filtering

- by default the event lists are filtered, and the filtered events are removed
- the filter expression can be controlled by the user
  - flagfilteredevents == true:
    In this case all events will be retained, and a flag column will be set to indicate what events would have been removed.

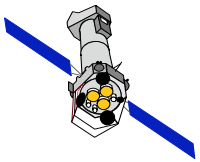
Fundamental final product of epicproc is the event list:

*[Imaging-Timing]Evts.ds (eg. 0193_0112570601_EMOS1_S001_ImagingEvts.ds)

+ BadPixel tables produced by (em)badpixfind

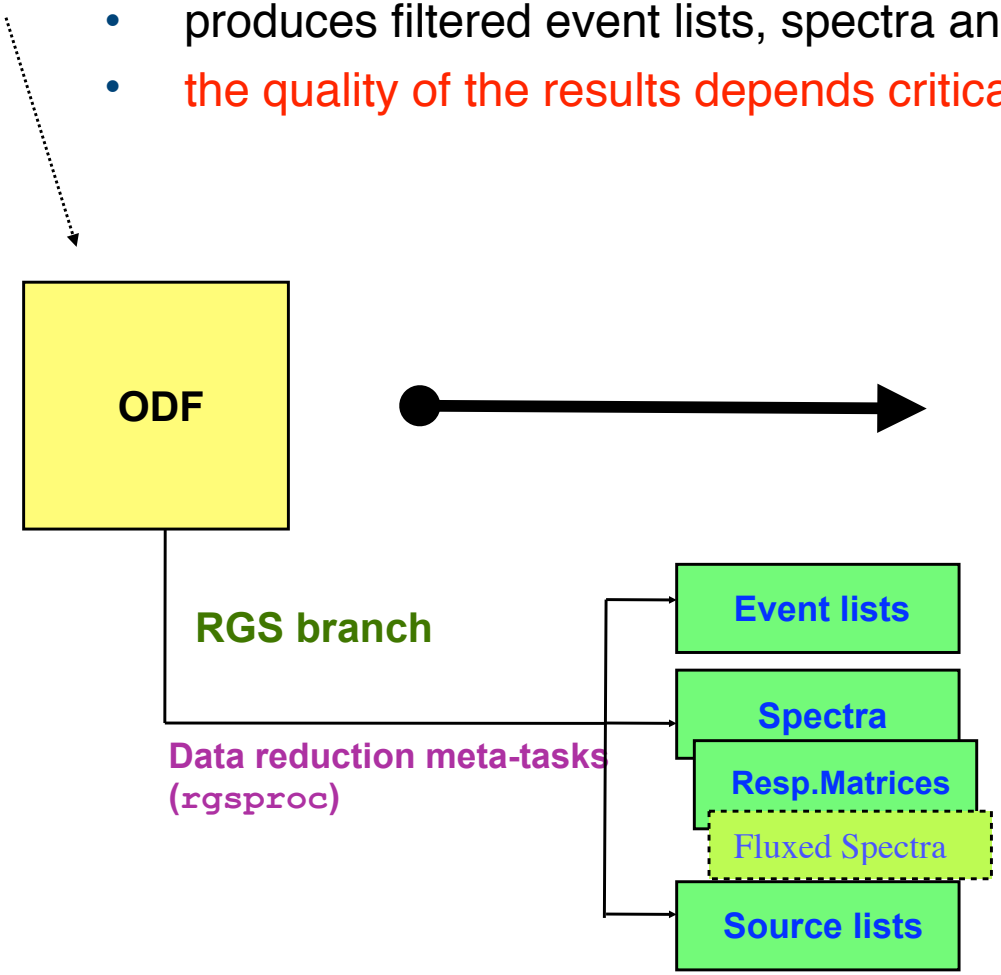
>> further scientific analysis based on work with event files [see Matteo's talk]

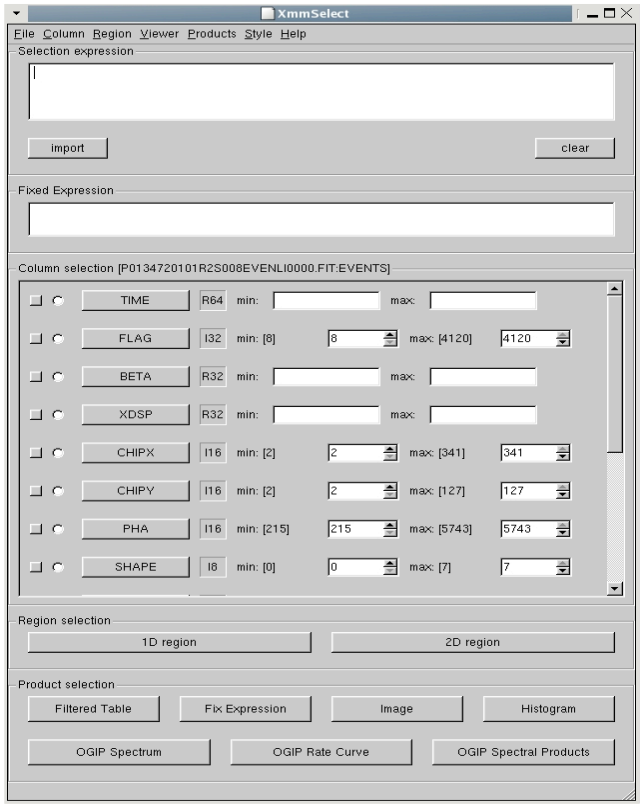
rgsproc



- meta-task: interface to 17 SAS tasks (that can also be run separately)
- controlled by ≈ 80 parameter switches
- five entry and final points (“processing stages”) >>
- produces filtered event lists, spectra and matrices
- the quality of the results depends critically on the source coordinates

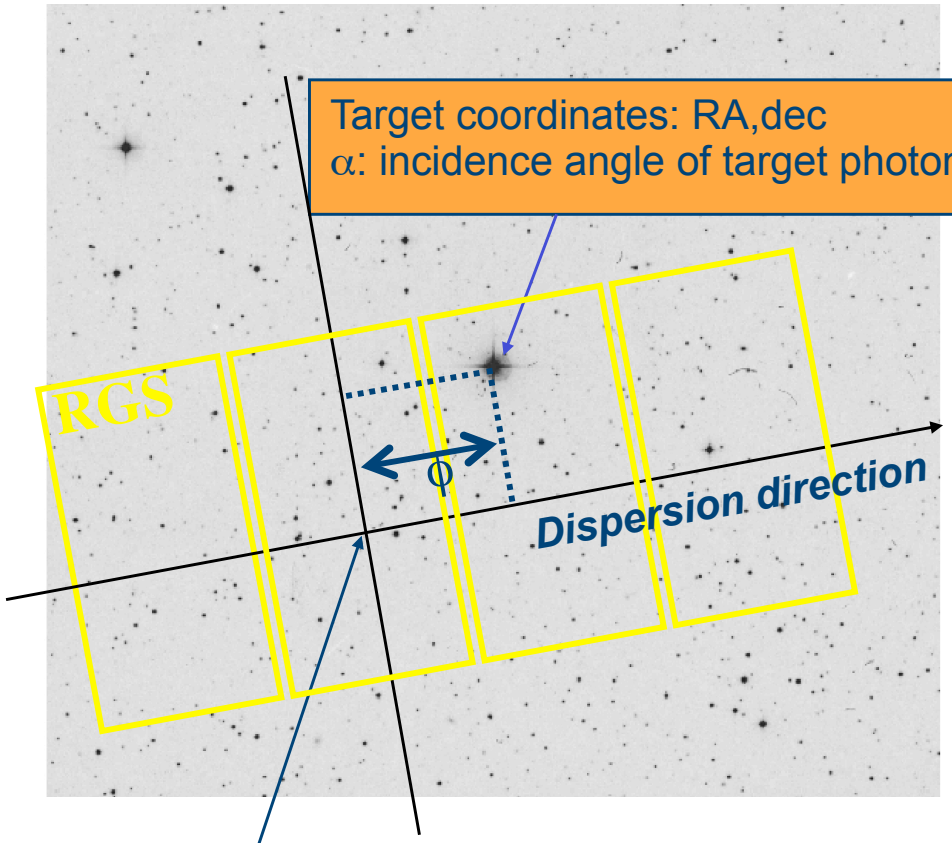
events
angles
filter
spectra
fluxing





Simplified scheme of the RGS FOV

why the coordinates are so important!:



According to the grating equation

$$\lambda = (\cos \beta - \cos \alpha)\, d / m$$

being $\alpha = \alpha_o + \phi\, F/L$

ϕ= f (RA- RA$_o$,dec-dec$_o$,P.A.)

⇩

the wavelength scale and the effective area depend on the position of the source in the FOV

1 arcsec ≈ 2.3 mÅ (45 km/s at 15 Å)

rgsproc: what does it do?

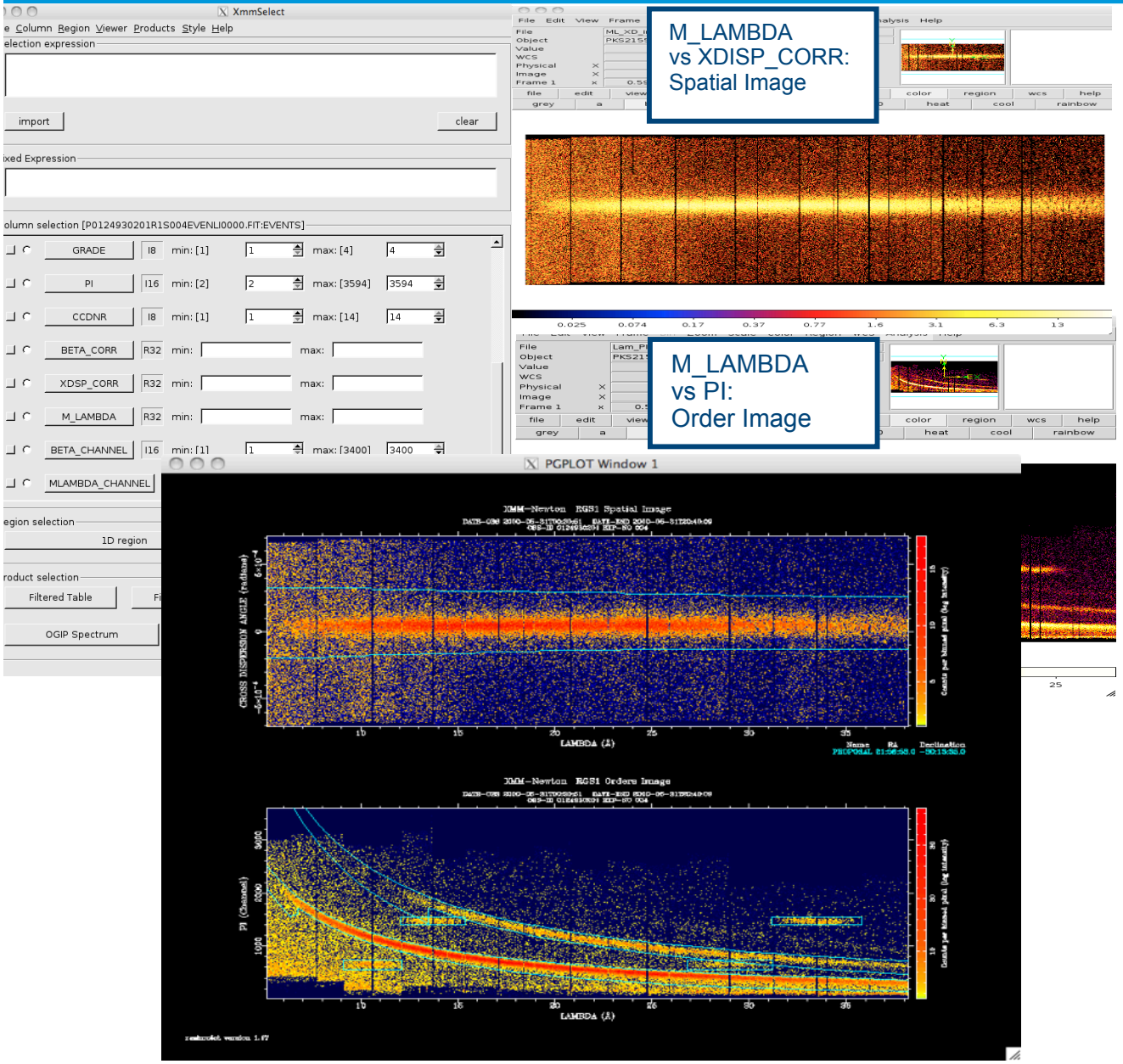
Stage	Task	Purpose	Output
Events	atthkgen	generates attitude file	Source list + intermediate combined event list
	attfilter	filters the attitude file	
	hkgtigen	generates housekeeping GTIs	
	rgsoffsetcalc	uses the diagnostic mode data for offset calculation	
	rgssources	creates the list of sources to processed	
	rgsframes	flags bad frames, convert RAW[XY] to readout node reference system ([XY]CORR), creates GTI for telemetry drops, calculates dead time	
	rgsenergy	performs energy calibrations, i.e. creates the PI column	
	rgsbadpix	flags bad pixels (CCF known + own analysis)	
	rgsevents	event reconstruction: total energy (ENERGY), "pattern" (GRADE/SHAPE), coordinates (CHIP[XY],BETA,XDSP)	
	evlistcomb	event list concatenation	

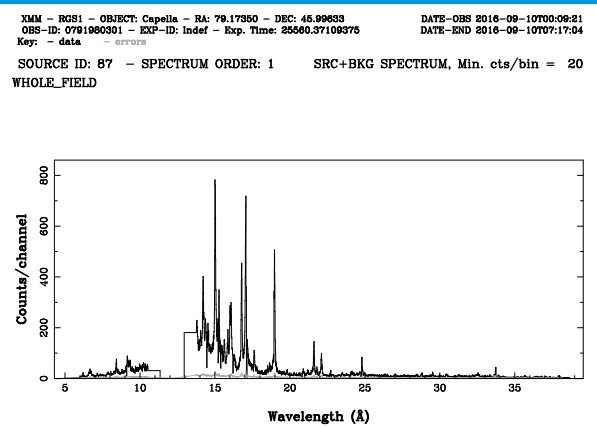
Source independent

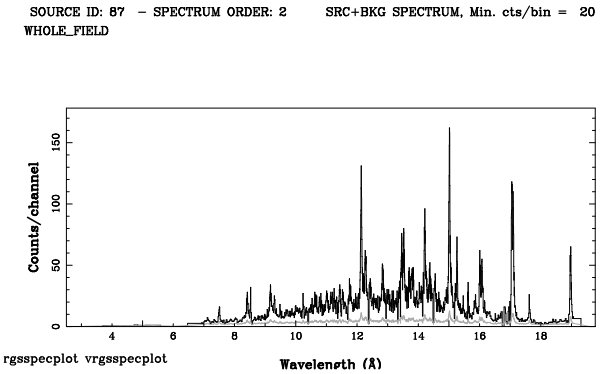
Angles	rgsangles	aspect correction (M_LAMBDA, XDSP_CORR)	Aspect correction
Filter	rgsfilter	creates filtered event list, removing unwanted frames and events and adding exposure maps	Final event list
Spectra	rgsregions	computes background and source extraction regions for each source	Source and background spectra
	rgsspectrum	extracts source and background spectra	
	rgsbkgmodel	generates model background (optional)	
Fluxing	rgsrmfgen	creates a response matrix	Response matrices and combined spectrum in physical units
	rgsfluxer	combines a collection of RGS spectra into one "fluxed" spectrum	

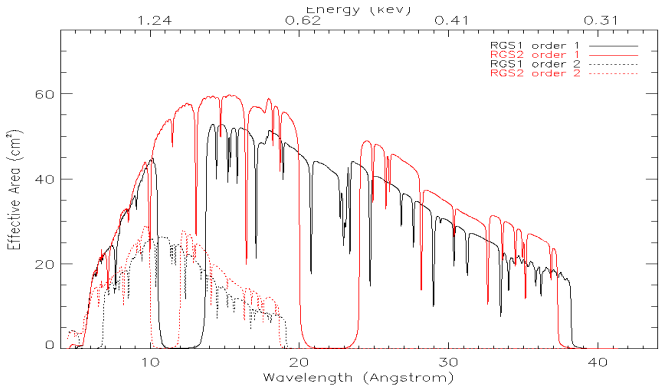
Source dependent

rgsproc does a lot









What do I get after processing?

For each RGS and exposure:

File	Content	rgsproc (default)	PPS
P0123456701R1S004EVENLI0000.FIT	Filtered Event List	Y	Y
P0123456701R1S004SRCLI_0000.FIT	Source List	Y (coord from proposal)	Y (coord from EPIC src list)
P0123456701R1S004BGSPEC1001.FIT P0123456701R1S004BGSPEC2001.FIT	Background Spectra (1st and 2nd order)	Y	Y
P0123456701R1S004SRSPEC1001.FIT P0123456701R1S004SRSPEC2001.FIT	Source Spectra (1st and 2nd order)	N	Y
P0123456701R1S004SBSPEC1001.FIT P0123456701R1S004SBSPEC2001.FIT	Source+Bkg Spectra (1st and 2nd order)	Y	Y
P0123456701R1S004RSPMAT1001.FIT P0123456701R1S004RSPMAT2001.FIT	Response Matrices (1st and 2nd order)	Y	only for 1st order
P0123456701R1S004IMAGE_0000.FIT P0123456701R1S004IMAGE_0000.PNG	Dispersion-CrossDispersion Image	N	Y
P0123456701R1S004ORDIMG0000.FIT P0123456701R1S004ORDIMG0000.PNG	Dispersion-Energy Image	N	Y
P0123456701R1S004EXPMAP0000.FIT	Exposure Map	Y	Y
P0123456701R1S004FBKTSR0000.FIT	Flare Background Timeseries	N	Y
P0123456701R1S004SRSPEC0001.PDF	Source Spectra (PDF)	N	Y

For each observation:

File	Content	rgsproc (default)	PPS
P0123456701OBX000fluxed1000.FIT P0123456701OBX000fluxed2000.FIT	Source Fluxed Spectra (1st and 2nd order)	Y	Y

Getting started with SAS I: the ODF

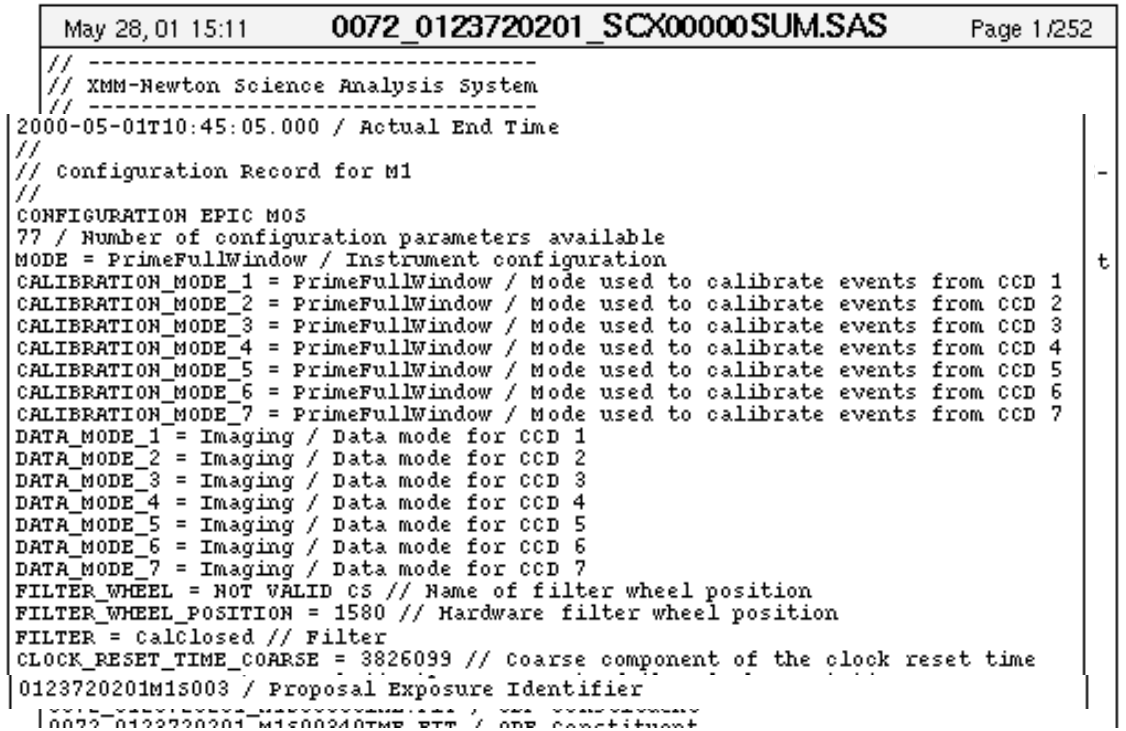
SAS needs for processing the ODF an advanced Summary File, produced by odfingest, extending an original summary file with data extracted from HK + calibration files

odfingest operates on the ODF directory $SAS_ODF

setenv SAS_ODF <odf_dir>

Command: > odfingest odfdir=$SAS_ODF

will produce a Summary file <SUM> = RRRR_OOOOOOOOOO_SCX00000SUM.SAS in the working directory



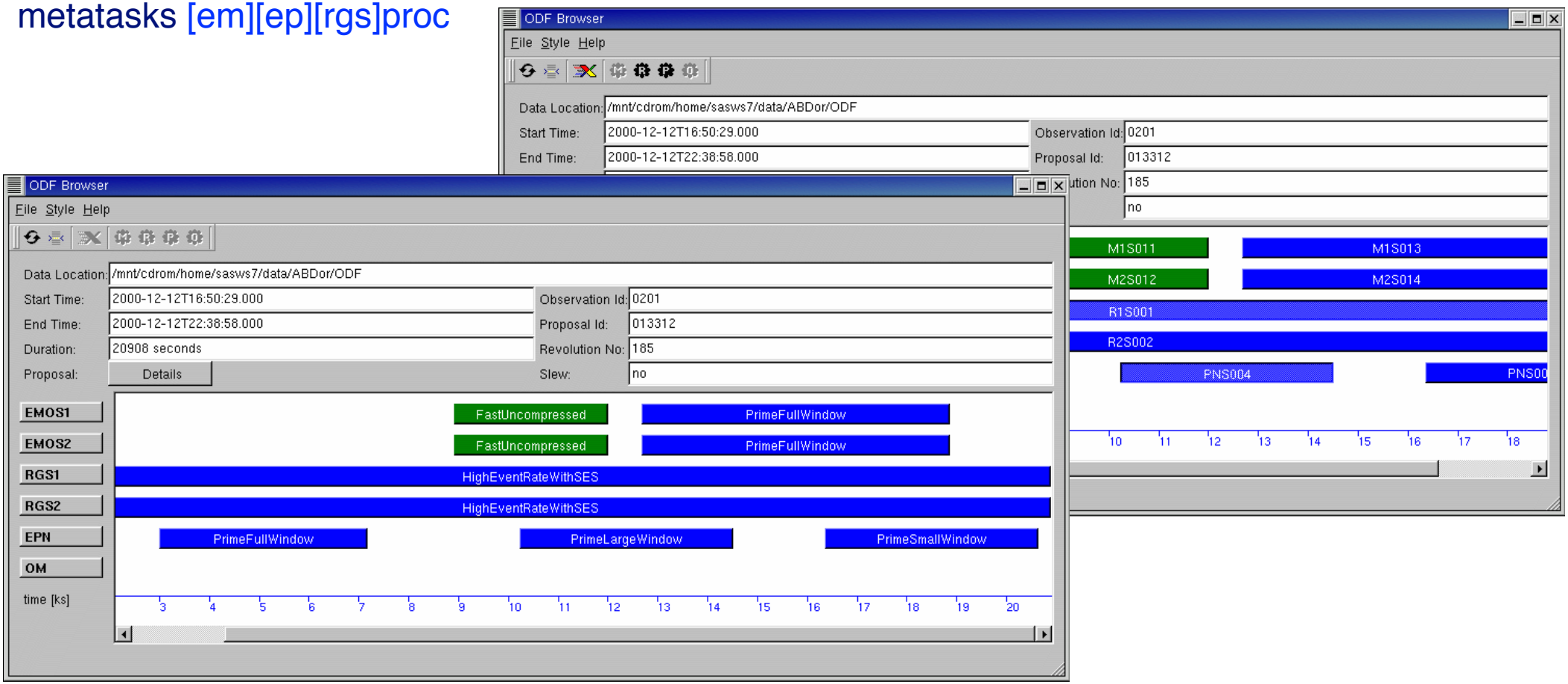
After the Summary file has been produced:

setenv SAS_ODF <SUM>

odfbrowser

A task to view the contents of the ODF and more …

An ODF is constituted in the rule by hundreds of files. odfbrowser displays graphical summaries of an observation, allowing the user to select any number of exposures and launch the metatasks [em][ep][rgs]proc



Getting started with SAS II: the CCF

XMM-Newton calibration data is contained in Current Calibration File (CCF)

CCF = the collection of all the XMM-Newton calibration files ever made public (= hundreds!)
Note: the calibration files are updated continuously >>> NO CCF version number

Calibration Index File (CIF) necessary for data analysis, pointing to the relevant files, according to:

- observation date
- analysis date

cifbuild operates on the calibration directory $SAS_CCFPATH setenv SAS_CCFPATH <ccf_dir>

Command: > cifbuild
produces a FITS file ccf.cif with extension CALINDEX >

After the Calibration Index file has been produced:
setenv SAS_CCF ccf.cif

File Edit Tools

	TELESCOP 4A	SCOPE 6A	TYPEID 32A	ISSUE I	VALDATE 19A yyyy:dd:mmZhh:mm:ss
40	XMM	EPN	LINCOORD	9	1998-01-01T00:00:00
41	XMM	EPN	MODEPARAM	3	1999-01-01T00:00:00
42	XMM	EPN	PATTERNLIB	1	1998-01-01T00:00:00
43	XMM	EPN	QUANTUMEF	8	2000-01-01T00:00:00
44	XMM	EPN	REDIST	5	1998-01-01T00:00:00
45	XMM	EPN	TIMECORR	4	1998-01-01T00:00:00
46	XMM	OM	ASTROMET	8	1998-01-01T00:00:00
47	XMM	OM	BADPIX	2	1998-01-01T00:00:00
48	XMM	OM	COLORTRANS	5	1998-01-01T00:00:00
49	XMM	OM	DARKFRAME	3	1998-01-01T00:00:00
50	XMM	OM	DIFFUSEGALA	1	1998-01-01T00:00:00
51	XMM	OM	HKPARMINT	3	1999-01-01T00:00:00
52	XMM	OM	LARGESCALESENS	2	1998-01-01T00:00:00
53	XMM	OM	LINCOORD	1	1998-01-01T00:00:00
54	XMM	OM	PHOTTONAT	3	1998-01-01T00:00:00
55	XMM	OM	PIXTOPIXSENS	3	1998-01-01T00:00:00
56	XMM	OM	PSF1DRB	4	1998-01-01T00:00:00
57	XMM	OM	QUICKMAG	2	1998-01-01T00:00:00
58	XMM	OM	ZODIACAL	1	1998-01-01T00:00:00
59	XMM	RGS1	ADUCONV	5	2000-02-06T16:49:60
60	XMM	RGS1	BACKGROUND	1	1998-01-01T00:00:00
61	XMM	RGS1	BADPIX	5	2000-02-06T16:49:60
62	XMM	RGS1	CALSOURCEDATA	1	1998-01-01T00:00:00
63	XMM	RGS1	CLOCKPATTERNS	1	1998-01-01T00:00:00
64	XMM	RGS1	CROSSPSF	2	2000-01-01T00:00:00
65	XMM	RGS1	CTI	2	2000-02-06T16:49:60
66	XMM	RGS1	DARKFRAME	4	1998-01-01T00:00:00
67	XMM	RGS1	HKPARMINT	6	1999-01-01T00:00:00
68	XMM	RGS1	LINCOORD	7	1998-01-01T00:00:00
69	XMM	RGS1	LINESPREADFUNC	3	1999-01-01T00:00:00

NOTICE: any file with extension CALINDEX is valid as SAS_CCF

cifbuild

cifbuild uses single CCF keywords:

- VALDATE as start of calibration validity period
- EVALDATE as end of validity period
- DATE as analysis validity period

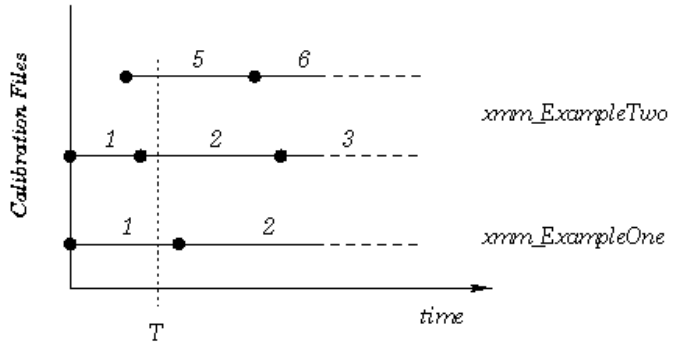


Figure 2: Current calibration file with two files: update. At the time T the current calibration file consists of *xmm_ExampleOne_0001.ccf* and *xmm_ExampleTwo_0005.ccf*

Rule: out of all the CCF calibration files take the highest issue with VALDATE lower AND EVALDATE higher than observation date AND DATE lower than analysis date.

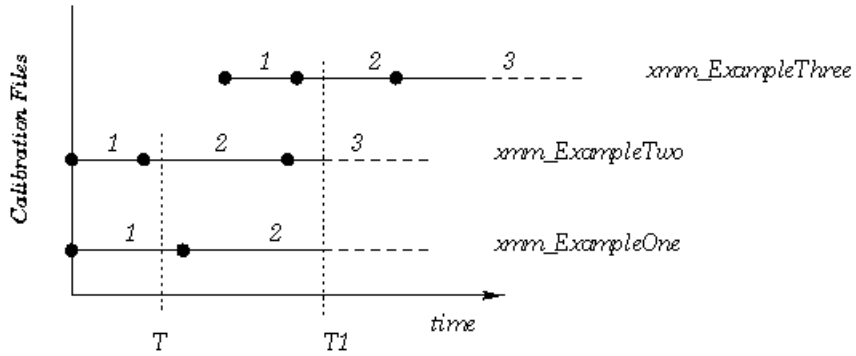


Figure 3: Current calibration file with three files. At the time $T1$ the current calibration file consists of *xmm_ExampleOne_0002.ccf* and *xmm_ExampleTwo_0009.ccf* and *xmm_ExampleThree_0002.ccf*

CIF / CCF on the web

- On the XMM-Newton calibration web pages:
  - updated cif can be generated on-line

    and compared to the one you generate
  - required (missing) CCF constituents can also be downloaded
  - local CCF library can be mirrored from XMM web site via the rsync or mirror commands (see doc web pages)
  - there is a “valid” CCF library (1.1 GB instead of > 6 GB)

    (if you don’t have it yet you can get it from me...)

- CCF release notes shall be consulted, at least periodically.
  - Subscribing to the CCF mailing list is also useful, to get the RNs only when there is something new.

Getting started with SAS III: the basic steps

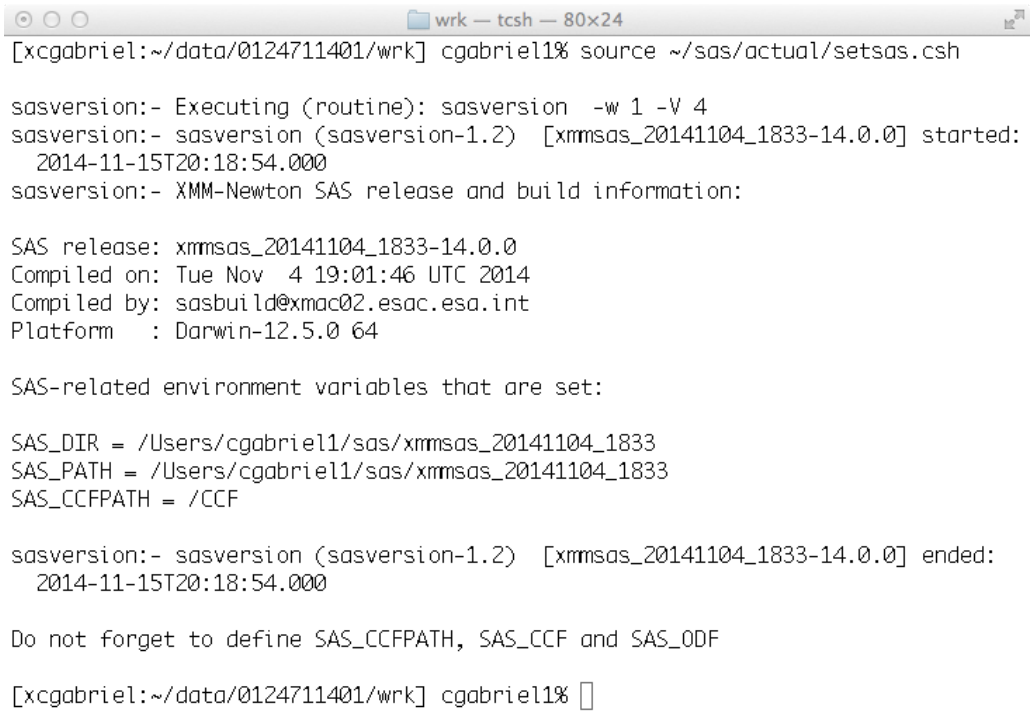
Summary of basic steps to start using SAS:
(csh / tcsh version)

source <SAS-DIR>**/setsas.csh**

setenv SAS_ODF <ODF-DIR>
setenv SAS_CCFPATH <Calibration Files-DIR>
cifbuild
setenv SAS_CCF ccf.cif

odfingest
set sumfile=`ls -1 *SUM.SAS`
setenv SAS_ODF $sumfile

>> ready to start working on the ODF data
located in <ODF-DIR>



Getting started with SAS III: the basic steps

OR, if you use the bash / sh / ksh:

. <SAS-DIR>/setsas.csh

export SAS_ODF =<ODF-DIR>

export SAS_CCFPATH =<Calibration Files-DIR>

cifbuild

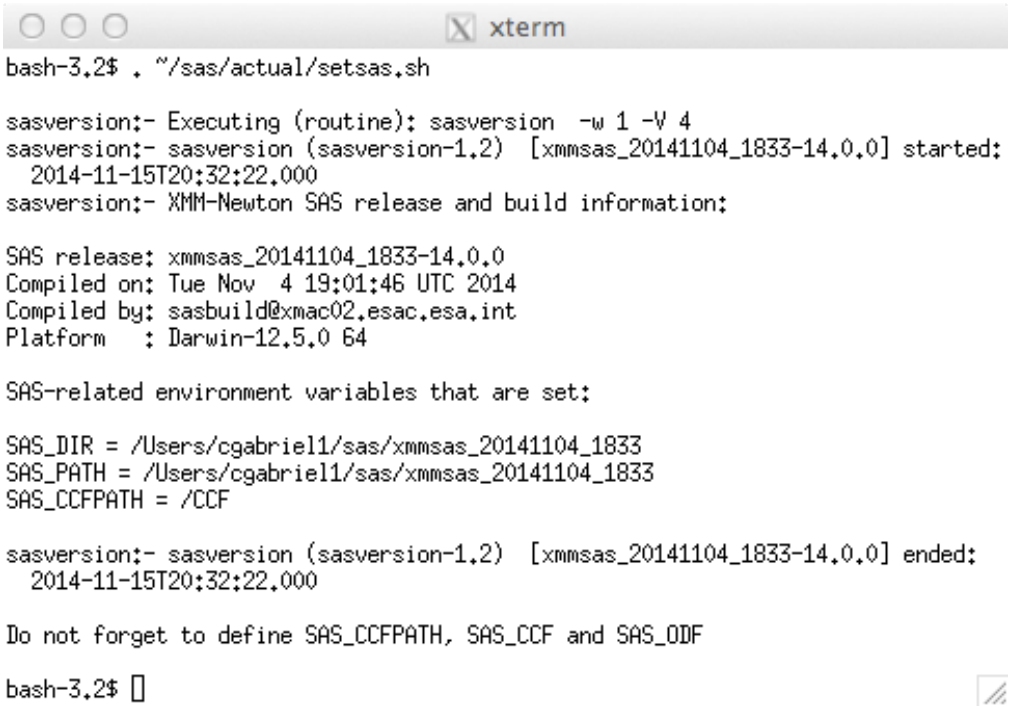
export SAS_CCF=ccf.cif

odfingest

export sumfile=`ls -1 *SUM.SAS`

export SAS_ODF=$sumfile

>> ready to start working on the ODF data
located in <ODF-DIR>



You can locate all these commands into (c)sh command files

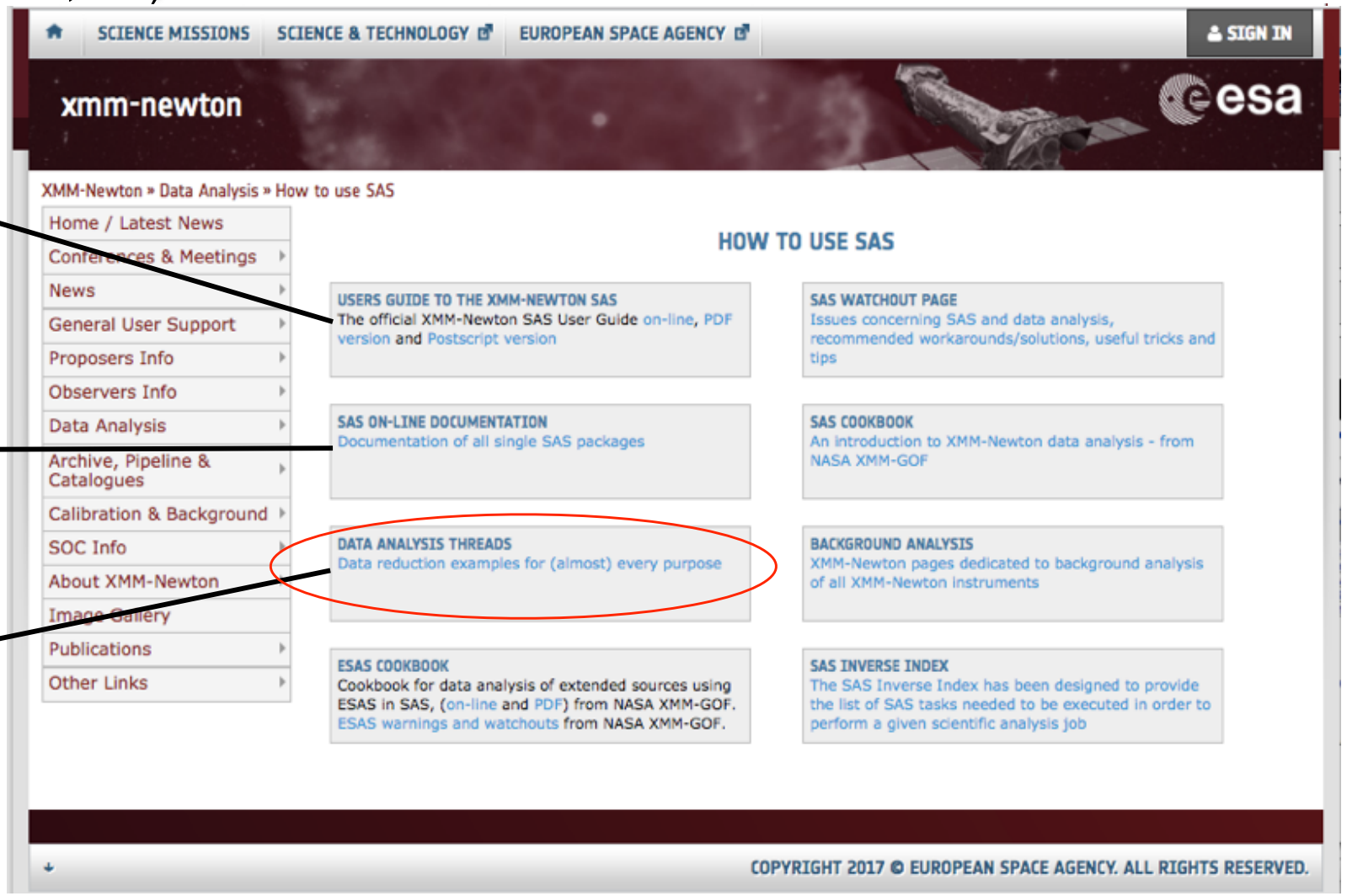
Getting started with SAS IV: all the information

- SAS public web page: https://www.cosmos.esa.int/web/xmm-newton/sas (download, installation, information, etc)
- ...How to use SAS

General guide

Task by task info

Analysis recipes



The SAS threads

XMM-Newton » Data Analysis » How to use SAS » Data Analysis Threads

- Home / Latest News
- Conferences & Meetings
- News
- Genera
- Propos
- Observ
- Data A
- Archive Catalog

SAS THREADS

EPIC RELATED THREADS

All in one go: from raw data (ODF) to science products		
- Analysis chain for point-like sources: `xmmextractor`	command line	
Step-by-Step		
Event list generation:		
- How to reprocess ODFs to generate calibrated and concatenated EPIC event lists	command line	
Filtering against high background:		
- How to filter EPIC event lists for flaring particle background	command line & GUI version	
Light curve generation:		
- Extraction of a light curve for a point-like source (EPIC and RGS)	command line	GUI version
Spectrum extraction:		
- Extraction of MOS spectra from point-like sources	command line	GUI version
- Extraction of MOS spectra from point-like sources taken in timing mode	command line	
- Extraction of pn spectra from point-like sources	command line	GUI version
- Extraction of pn spectra from point-like sources taken in timing mode	command line	
- Extraction of spectra in a few clicks: `especget`		GUI version
- Combining the spectra of the 3 EPIC cameras	command line	
- Overlapping EPIC data treatment: `multixmmselect`		GUI version
Point Spread Function (PSF) generation:		
- 2-D PSF à la carte	command line	
More complex analysis for bright sources		
- Dealing with EPIC Out-of-Time (OoT) events	command line	
- How to evaluate and test pile-up in an EPIC source	command line	
Handling of EPIC background		
- How to use EPIC instrumental background files	command line	
ESAS:		
- Creation of EPIC background subtracted, exposure corrected images	command line	
- Creation of EPIC merged background subtracted and exposure corrected images	command line	
- Creation of EPIC spectral analysis files for a cluster radial profile	command line	
Images:		
- A shell script to create attractive EPIC-pn & MOS combined images	dedicated Web page	
- How to Generate Vignetting-corrected Background-subtracted EPIC Images	command line	
Source detection		
- EPIC source finding thread in one go: `edetect_chain`	command line	
- EPIC source finding thread: step-by-step	command line	
- EPIC source finding in overlapping exposures	command line	
Slew data processing		
- How to process EPIC slew data	command line	

	command line	
	command line	
	command line	
	command line	

RGS RELATED THREADS

All in one go: from raw data (ODF) to science products		
- Analysis chain for point-like sources: `xmmextractor`	command line	
Step-by-Step		
- How to reduce RGS data and extract spectra of point-like sources	command line	
- `rgsproc`, coordinates and masks	command line	
Light curve generation:		
- Extraction of a light curve for a point-like source (EPIC and RGS)	command line	GUI version
More complex analysis for the very bright sources		
- Pile-up in the RGS: how to prevent it, evaluate its existence and make corrections	command line	

OM RELATED THREADS

All in one go: from raw data (ODF) to science products		
- Analysis chain for point-like sources: `xmmextractor`	command line	
Step-by-Step		
- OM image mode data processing chain	processing chain	command line
- OM fast mode data processing chain	processing chain	command line
- OM Grism processing chain	processing chain	command line
- Interactive OM photometry	command line	
- Converting OM data to OGIP II format (use in `xspec`)	command line	

A SAS thread

This thread contains a step-by-step recipe to extract light curves of a point-like source for all the X-ray cameras, subt

Expected Outcome

Corrected light curves of XMM-Newton EPIC and RGS instruments.

SAS Tasks to be Used

- evselect
- epiclccorr
- rgslccorr
- barycen

Prerequisites

- SAS Start-up Thread
- How to reprocess ODFs to generate calibrated and concatenated EPIC event lists Thread
- How to reduce RGS data and extract spectra of point-like sources Thread

Useful Links

- How to evaluate the pile-up fraction thread

Caveats

Last Reviewed: ***29 May 2013, for SAS v13.0***

Last Updated: ***29 May 2013***

Procedure

EPIC

As an example case, we will consider the extraction of a light curve from a pn event list (PN_evt.fits). The same re

1. Set up your SAS environment (see **Prerequisites** for this thread at the top of the page).
2. Be aware: if you are interested in very short time periods, such as they appear in pulsars or cataclysmic variabl arrival time of a photon is shifted as is it would have been detected at the barycentre of the solar system (the c the data are comparable. The SAS task barycen performs this correction. As barycen overwrites the TIME colum

```
cp PN_evt.fits PN_evt_barc
barycen table=PN_evt_barce
```

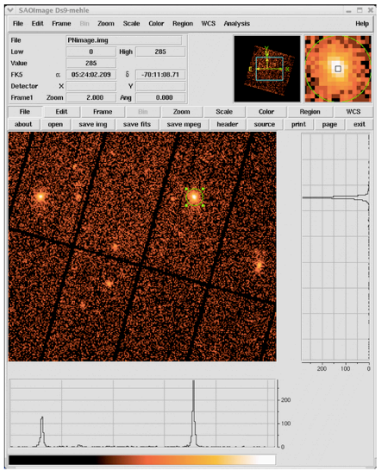
3. Extract an image (in sky coordir

```
evselect table=PN_evt.fits
  xcolumn=X ycolumn=Y xim
```

4. Display the image

```
imgdisplay withimagefile=t
```

5. Select the region, from which th



6. Double-click with the cursor on the defined region. A window pops up, showing the properties of the region (Fig.2). Write down the coordinates of the Centre (25910.5, 25870.5) and the Radius(400).

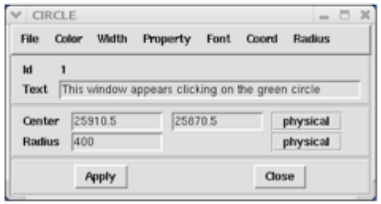


Fig.2: Selection region properties window, popped-up by double-clicking on the region in the main ds9 window

Units of sky coordinates (X,Y) are 0.05 arcsec, hence the radius in our example is 20 arcsec.

7. Now you can extract a source+background light curve, using all the selection region and including a quality selection appropriate for a light curve extraction. For PN, taking good events, singles and doubles with an energy range between 200 and 10000 eV (#XMMEA_EP && (PATTERN<=4) && (PI in [200:10000])). For MOS, taking good events, singles, doubles, triples and quadruples with an energy range between 200 and 10000 eV (#XMMEA_EM && (PATTERN<=12) && (PI in [200:10000])). In the example, the bin size is 100 seconds.

```
evselect table=PN_evt.fits energycolumn=PI expression='#XMMEA_EP&&(PATTERN<=4)&& \
  ((X,Y) IN circle(25910.5,25870.5,400))&&(PI in [200:10000])' \
  withrateset=yes rateset="PN_source_lightcurve_raw.lc" timebinsize=100 \
  maketimecolumn=yes makeratecolumn=yes
```

The parameter makeratecolumn=yes produces a light curve in count rates (with errors). Otherwise the light curve is produced in counts (with errors).

8. Repeat steps #4 to #6 above to determine the region, from which the background light curve is to be extracted. We will assume in the following that the extraction region corresponds to an annulus, centered in (25910.5,25870.5) and with inner and outer radii 1000 and 2000 pixels, respectively.
9. Extract a background light curve, using all the selection expressions defined so far, and the same bin size (100 seconds) and energy range as for the source+background light curve

```
evselect table=PN_evt.fits energycolumn=PI expression='#XMMEA_EP&&(PATTERN<=4)&& \
  ((X,Y) IN annulus(25910.5,25870.5,1000,2000))&&(PI in [200:10000])' withrateset=yes \
  rateset="PN_light_curve_background_raw.lc" timebinsize=100 \
  maketime
```

The light curv

10. However, ligh efficiency, as source and b both light cur

```
epiclccorr
  bkgtsli
```

11. Plot the resul

```
dsplot tab
```

This comman

11. Plot the resulting light curves, e.g.

```
dsplot table=PN_lccorr.lc withx=yes x=TIME withy=yes y=RATE
```

This command will launch the following xmgrace window

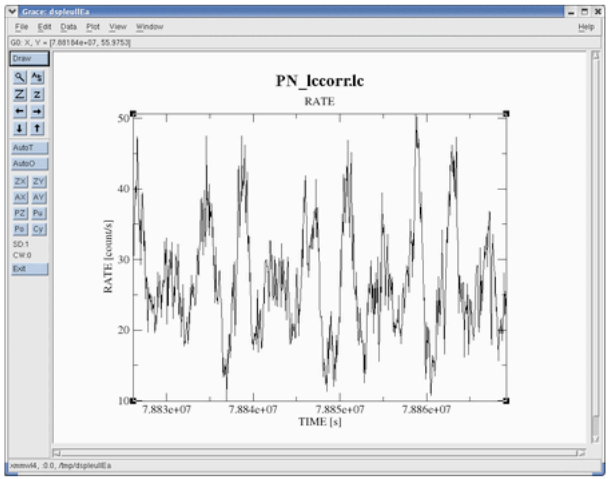
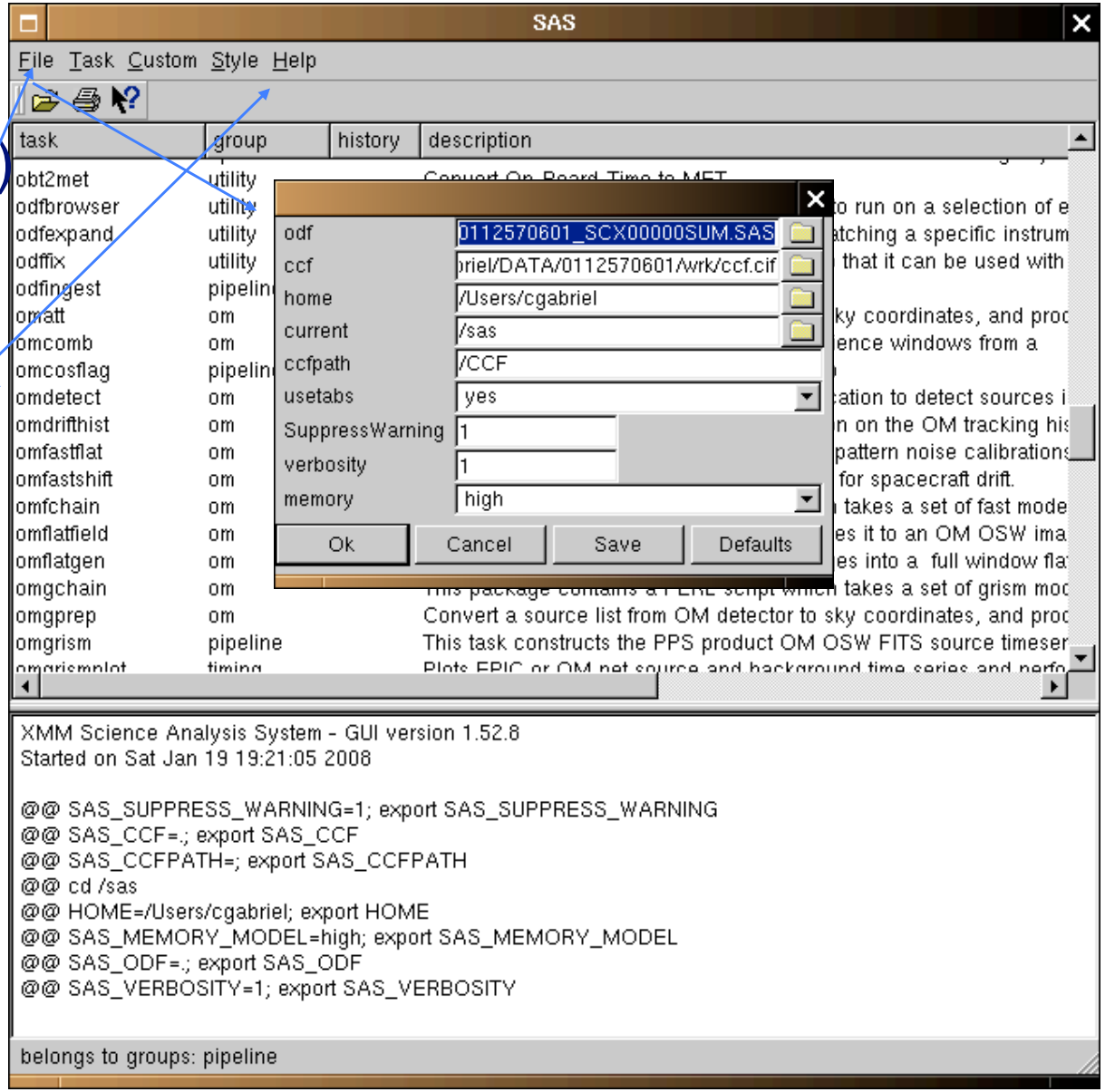


Fig.3: xmgrace window, containing the background-subtracted exposure-corrected light curve

The main GUI

- Access to all tasks (GUI call) and descriptions
- Setting general defaults
- Access to help pages
- Writing log into window and file (sas_log)

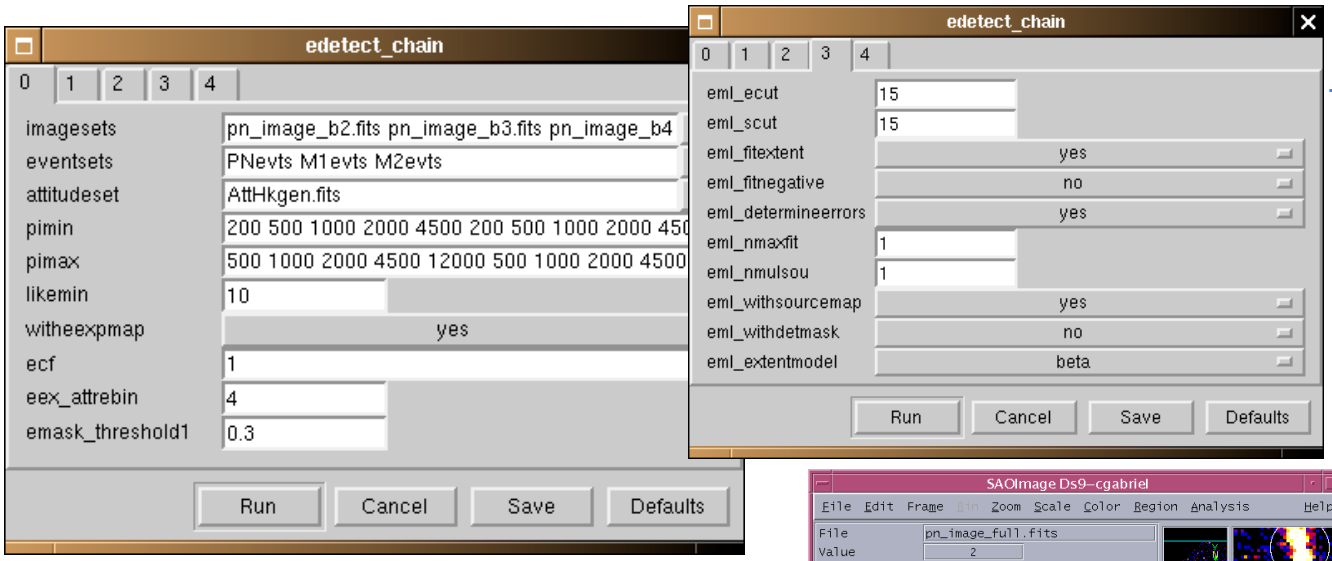


GUI or command line?

GUIs are very useful for beginners

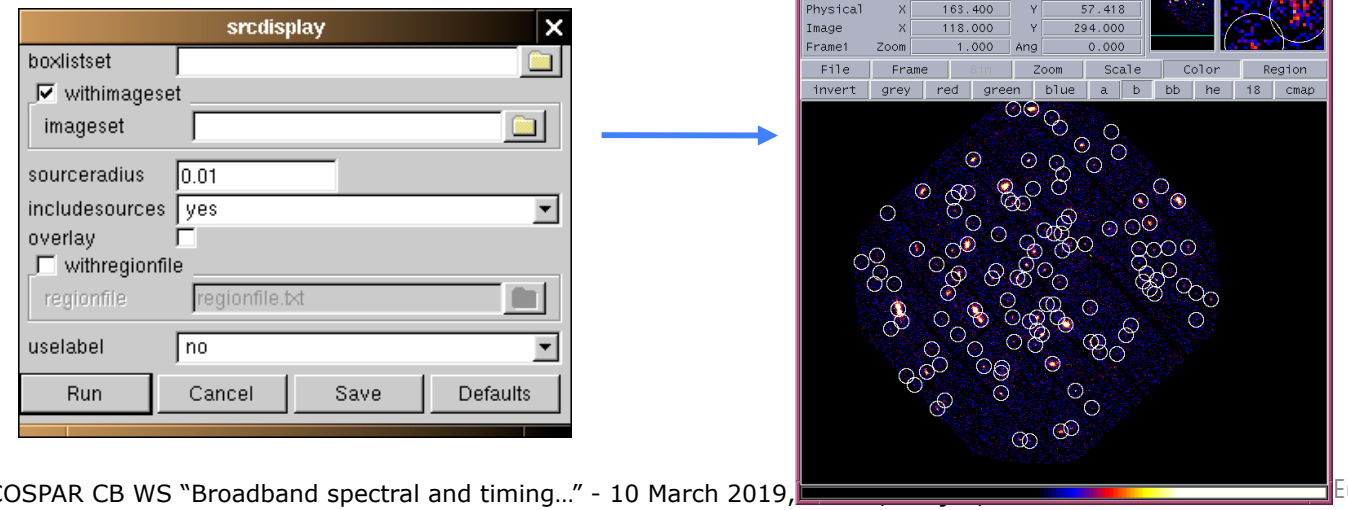
Every SAS task has its own GUI - they can be called by other (main) GUI or directly from the command line by > <task> -d:

edetect_chain –d



source list

srcdisplay -d



GUI or command line? BOTH

Line commands are, if you know them, faster in execution, and you learn how to concatenate them, so that ...

eventually, you can produce your own scripts

det_script - /Users/cgabriel/DATA/0112570601/wrk/srcdet/

File Edit Search Preferences Shell Macro Windows Help

```
evselect table=MOS1evt:EVENTS expression='#XMMEA_EM&&(PI>10000)&&(PATTERN==0)' rateset="m1_back_lightc.fits" timebinsize=10
withrateset=yes maketimecolumn=yes makeratecolumn=yes
evselect table=MOS2evt:EVENTS expression='#XMMEA_EM&&(PI>10000)&&(PATTERN==0)' rateset="m2_back_lightc.fits" timebinsize=10
withrateset=yes maketimecolumn=yes makeratecolumn=yes
evselect table=PNevt:EVENTS expression='#XMMEA_EP&&(PI>10000)&&(PATTERN==0)' rateset="pn_back_lightc.fits" timebinsize=10
withrateset=yes maketimecolumn=yes makeratecolumn=yes

dsplot table=m1_back_lightc.fits x=TIME y=RATE &
dsplot table=m2_back_lightc.fits x=TIME y=RATE &
dsplot table=pn_back_lightc.fits x=TIME y=RATE &

tabgtigen table=m1_back_lightc.fits expression="RATE < 0.35" gtiset=m1_back_gti.fits
tabgtigen table=m2_back_lightc.fits expression="RATE < 0.35" gtiset=m2_back_gti.fits
tabgtigen table=pn_back_lightc.fits expression="RATE < 1.00" gtiset=pn_back_gti.fits

evselect table=MOS1evt:EVENTS imagebinning='binSize' imageset='m1_image_full.fits' withimageset=yes xcolumn='X' ycolumn='Y'
ximagebinsize=40 yimagebinsize=40 expression='#XMMEA_EM&&(PI in [200:12000])&&(PATTERN in [0:12])&&(FLAG==0) &&
gti(m1_back_gti.fits,TIME)'
evselect table=MOS1evt:EVENTS imagebinning='binSize' imageset='m1_image_b1.fits' withimageset=yes xcolumn='X' ycolumn='Y'
ximagebinsize=40 yimagebinsize=40 expression='#XMMEA_EM&&(PI in [200:500])&&(PATTERN in [0:12])&&(FLAG==0) &&
gti(m1_back_gti.fits,TIME)'
evselect table=MOS1evt:EVENTS imagebinning='binSize' imageset='m1_image_b2.fits' withimageset=yes xcolumn='X' ycolumn='Y'
ximagebinsize=40 yimagebinsize=40 expression='#XMMEA_EM&&(PI in [500:1000])&&(PATTERN in [0:12])&&(FLAG==0) &&
gti(m1_back_gti.fits,TIME)'
evselect table=MOS1evt:EVENTS imagebinning='binSize' imageset='m1_image_b3.fits' withimageset=yes xcolumn='X' ycolumn='Y'
ximagebinsize=40 yimagebinsize=40 expression='#XMMEA_EM&&(PI in [1000:2000])&&(PATTERN in [0:12])&&(FLAG==0) &&
```

My answer: GUI & command line >> scripts

PPS or “proc” products? BOTH

All data → already reduced by PPS (SAS subset with default parameters)

Why reprocessing then (epproc, emproc, rgsproc, om?chain) ?

- old SAS/PPS used for that data
- newer (probably more accurate) calibration than the one used by PPS
- special needs for using non-default parameters

During 2012 general reprocessing (second time in 12 years)
>> XMM Science Archive repopulated with s/w and calibration as 2012 knowledge
+ source specific products: spectra and light curves
+ 3XMM - the largest X-ray sources catalogue

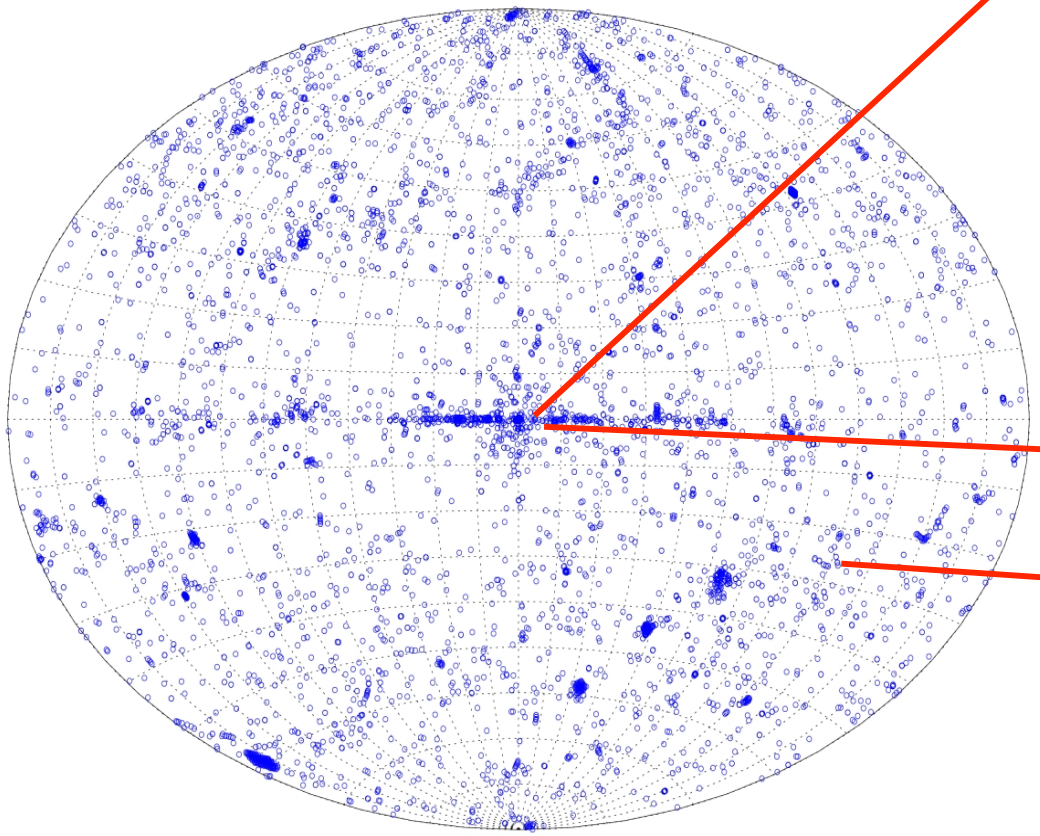
Bulk reprocessing will take place in 2019 … 4XMM

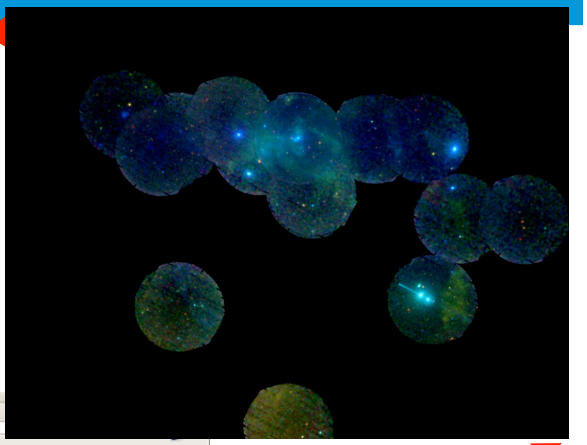
Even if you reduce yourself the data it is important to get a look into the PPS data beforehand, in ALL cases they will give a good impression about the contents

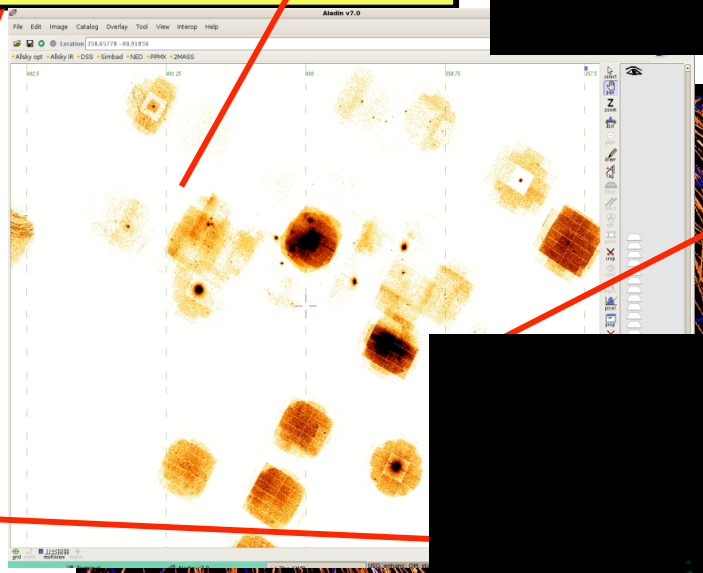
Where is the XMM-Newton data? In XSA

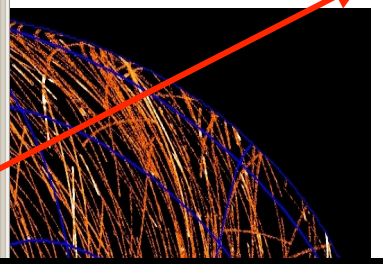
Contents of the XSA:

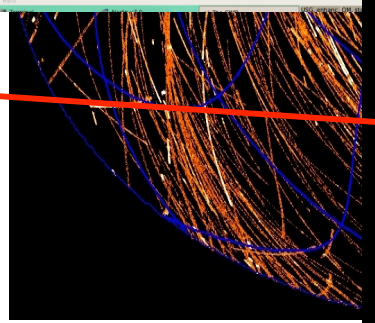
- ODF/PPS of ~ 14000 pointed observations
- SDF of ~ 4,000 slew observations
- SDF of ~ 200,000 Slew Survey sub-exposures
- 775.153 detections (3XMM-DR8 catalogue) / 531.454 unique sources
- 6.880.116 OM sources (XMMOM SUSS 3 catalogue)
- 72352 Slew Survey sources (XMMSL2)

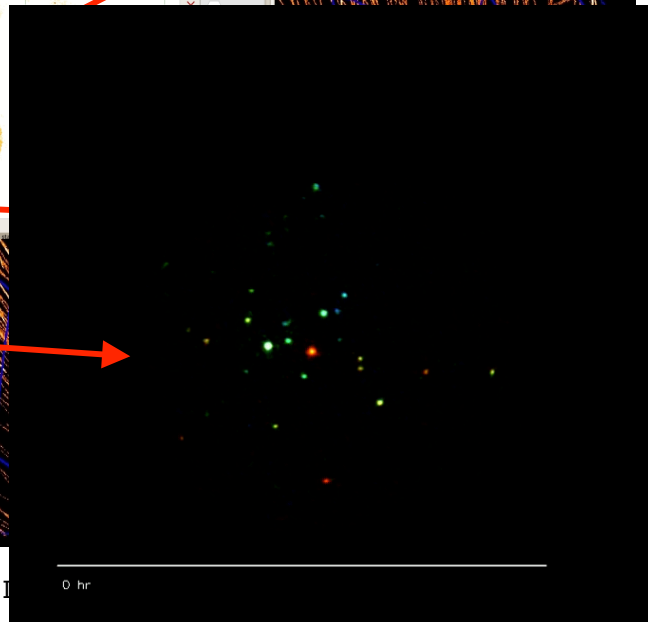




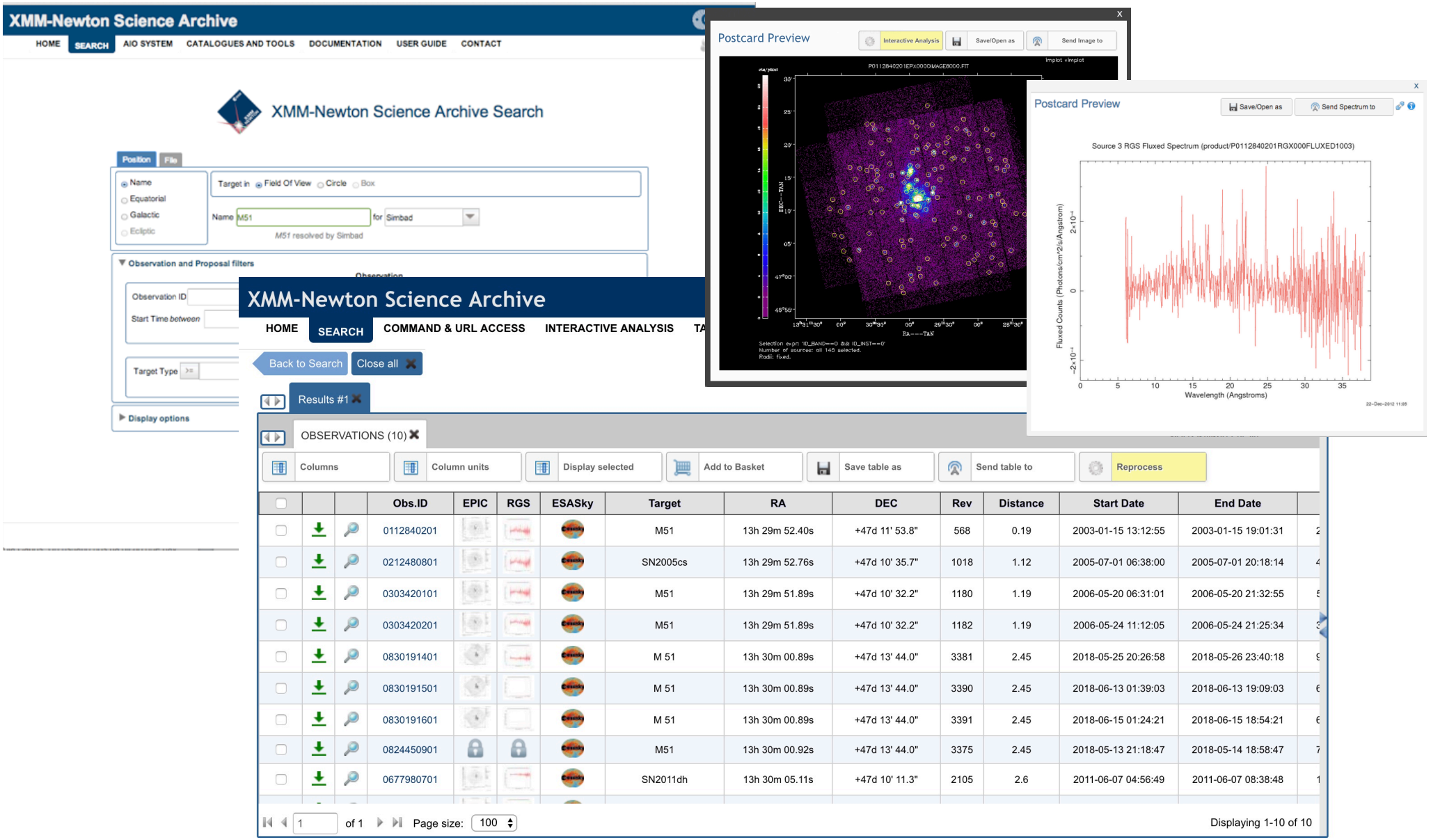




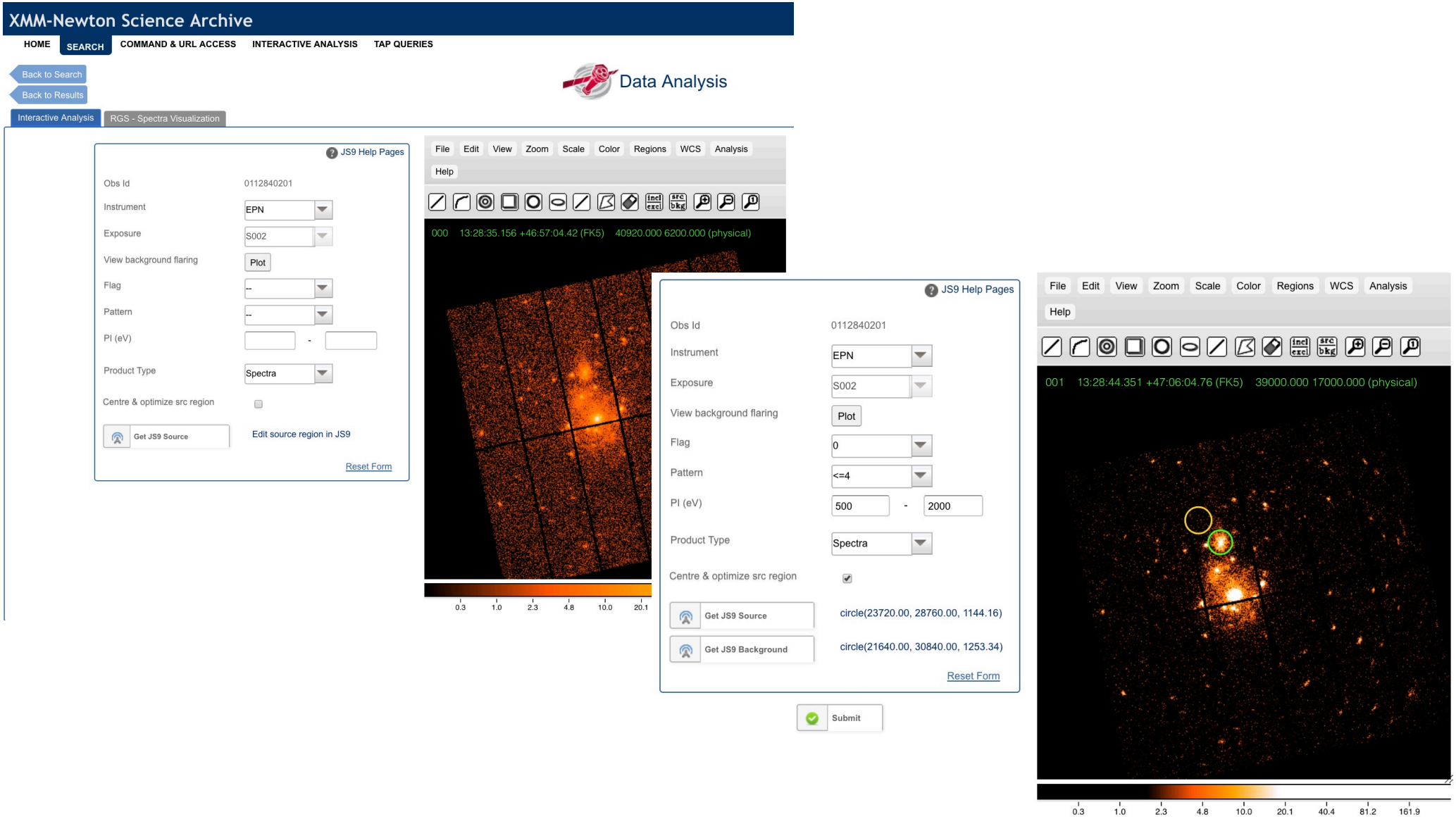




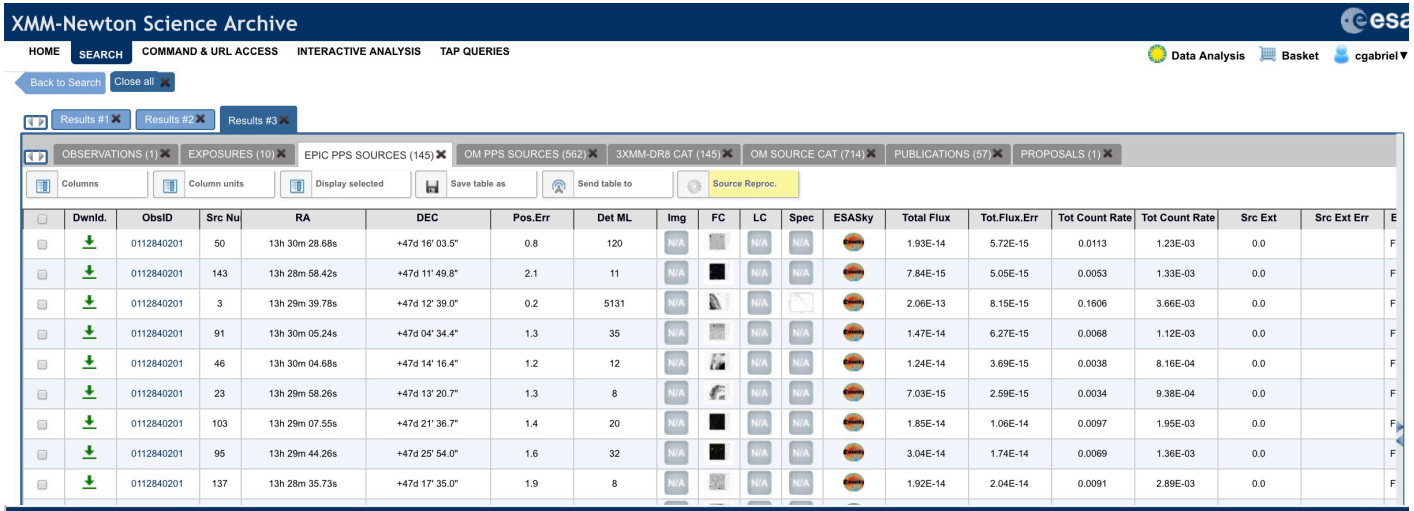
How to get XMM-Newton data? The XSA

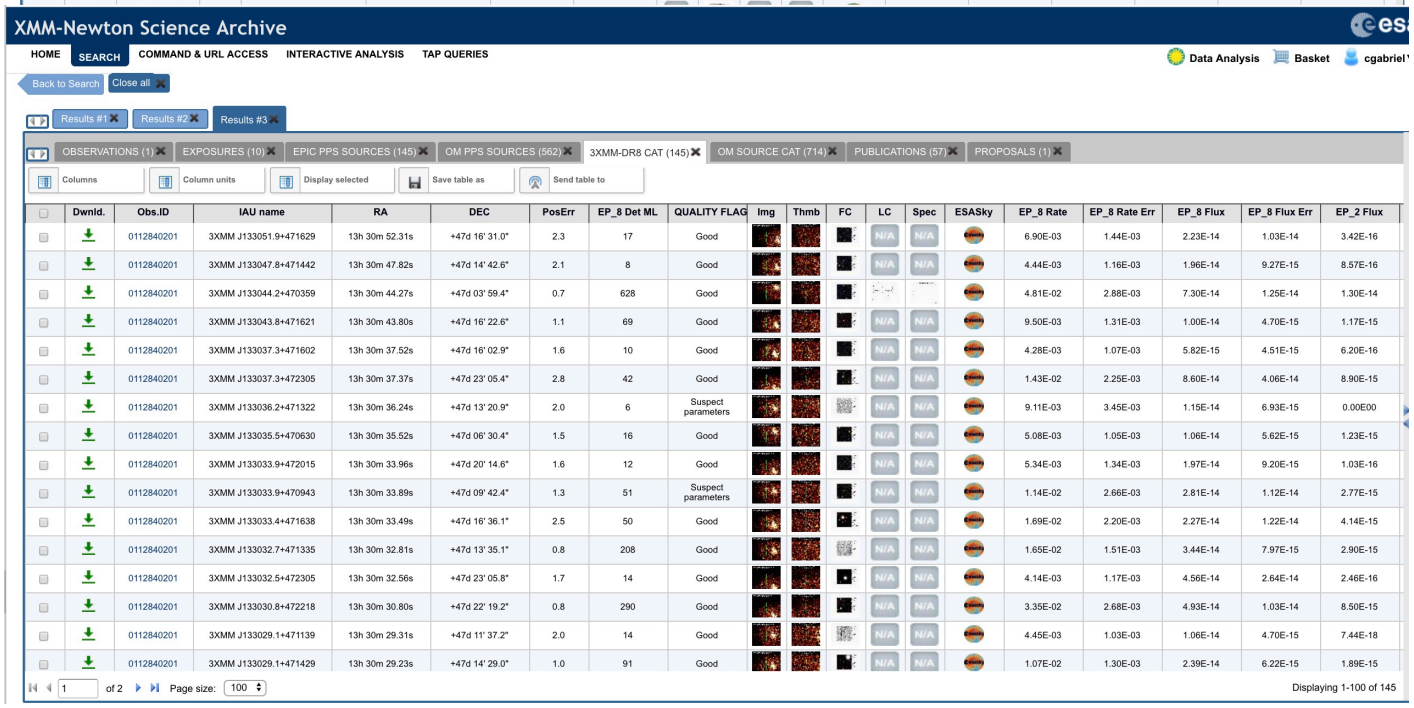


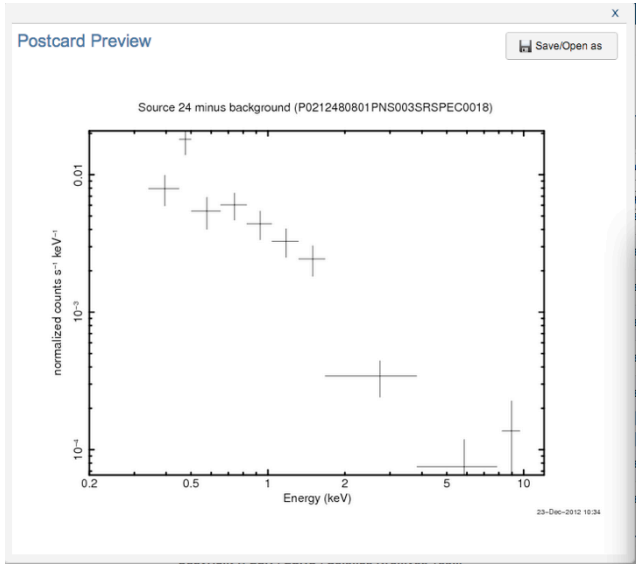
Processing from XSA (using RISA)

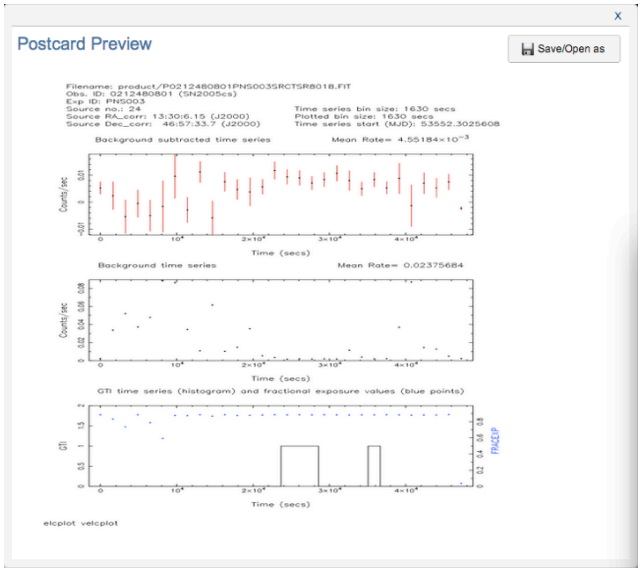


All the individual sources detected









How to find files in the XMM-Newton data forest

EPIC:

MIEVLI / *PIEVLI* files are the event list files in PPS
*[Imaging-Timing]*Evts.ds files are the corresponding ones in PROC
IMAGE_8000.FTZ are the FITS compressed whole camera images in PPS
The only EPIC products from e[m-p]proc are the event list files and Bad Pixel tables

RGS:

*R1*EVENLI* / *R2*EVENLI* are the event list files in PPS and PROC
*R1*SRCLI* / *R2*SRCLI* are the source list files in PPS and PROC
SRSPEC are the source subtracted spectra in PPS and PROC
Response matrices are products in PROC (*Matrix*), as well as fluxed spectra (*fluxed*)
Intermediate RGS products are only kept in PROC (not in PPS), therefore RGS analysis starting on intermediate point only possible with PROC products

OM:

Whole analysis done by pipeline – Check source detections (using eg. implot, ds9, ftools)
If necessary, re-start analysis at intermediate steps (omdetect in om[i-f]thread.html)

PPS Index Summary can help a lot to recognize correspondence FileType <-> FileName

How to find files in the XMM-Newton data forest

PPS Index Summary can help a lot to recognize correspondence FileType <-> FileName ... but also to find everything produced by PPS (and that's a lot!)

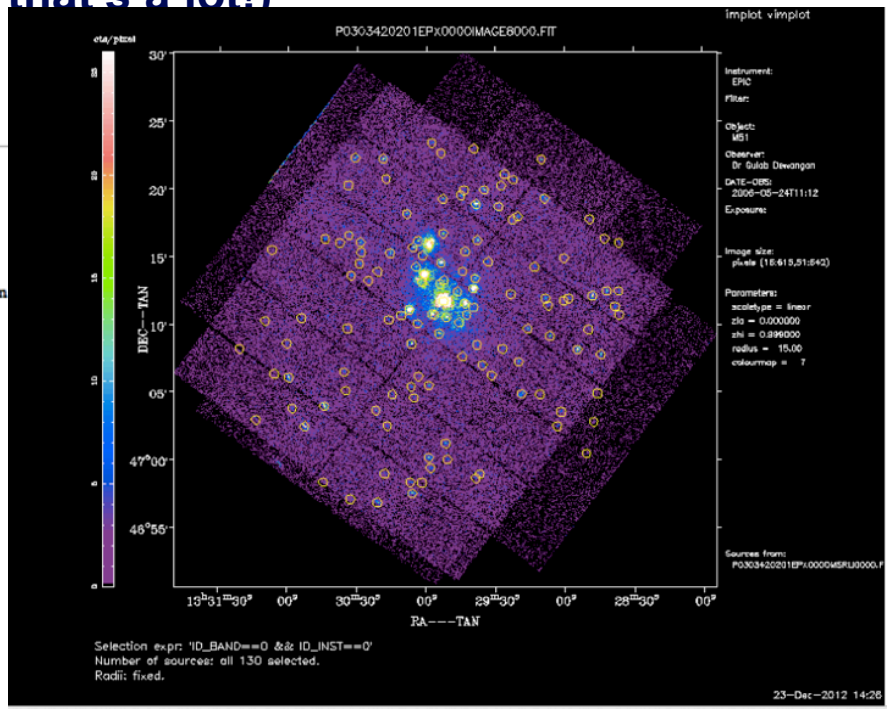
XMM SURVEY SCIENCE CENTRE esa

OBS Summary | PPS Summary | EPIC Summary | OM Summary | RGS Summary | Catalogue

0303420201 EPIC Processing Summary

Inst.	Exp. Id	Sched	Mode	Datamode	Filter	Position	Duration	Exposure	SrcDet	SSP	Flare Scrn
EPN	S003	Y	PrimeFullWindow	Imaging	THIN1		34997	22545	Y	Y	Y

Filename	Content	Band
P0303420201PNS003PIEVLI0000.FTZ	EPIC PN IMAGING MODE EVENT LIST	
P0303420201PNS003FBKTSR0000.FTZ	EPIC FLARE BACKGROUND TIMESERIES	
P0303420201PNS003FBKTSR0000.PDF	EPIC FLARE BACKGROUND TIMESERIES	
P0303420201PNS003EXPMAP1000.FTZ	EPIC EXPOSURE MAP	0.2 - 0.5
P0303420201PNS003IMAGE_1000.FTZ	EPIC IMAGE	0.2 - 0.5
P0303420201PNS003EXPMAP2000.FTZ	EPIC EXPOSURE MAP	0.5 - 1.0
P0303420201PNS003IMAGE_2000.FTZ	EPIC IMAGE	0.5 - 1.0
P0303420201PNS003EXPMAP3000.FTZ	EPIC EXPOSURE MAP	1.0 - 2.0
P0303420201PNS003IMAGE_3000.FTZ	EPIC IMAGE	1.0 - 2.0
P0303420201PNS003EXPMAP4000.FTZ	EPIC EXPOSURE MAP	2.0 - 4.5
P0303420201PNS003IMAGE_4000.FTZ	EPIC IMAGE	2.0 - 4.5
P0303420201PNS003EXPMAP5000.FTZ	EPIC EXPOSURE MAP	4.5 - 12.0
P0303420201PNS003IMAGE_5000.FTZ	EPIC IMAGE	4.5 - 12.0
P0303420201PNS003EXPMAP8000.FTZ	EPIC EXPOSURE MAP	0.2 - 12.0
P0303420201PNS003EXPMAP8000.PNG	EPIC EXPOSURE MAP	0.2 - 12.0
P0303420201PNS003IMAGE_8		
P0303420201PNS003IMAGE_8		



Sources processed by PP

Source	Inst.	Exp. Id	RA	Dec	RADec Err	Count	Band 8 Flux	EPIC Det ML	Inst. Det ML	Ontime	EPIC Extent	EPIC Flags	Srcdet
Dec (Hex)			[hms]	[° ′ ″]	[″]		[erg/s/cm^2]			[s]	[image pixels]		
1 (001)	EPN	S003	13 29 52.54	47 11 44.6	0.2	12549.8 ±138.9	1.5e-12 ±3.32e-14	50374	22575	22545	2 ±0	FFTFFFFFFFFF	Y

Filename	Content	Band
P0303420201PNS003SRSPEC0001.FTZ	EPIC SOURCE SPECTRUM	
P0303420201PNS003BGSPEC0001.FTZ	EPIC SOURCE BACKGROUND SPECTRUM	
P0303420201PNS003SPCPLT0001.PDF	EPIC SOURCE SPECTRUM PLOT	
P0303420201PNS003SRCTSR8001.FTZ	EPIC SOURCE TIMESERIES	0.2 - 12.0
P0303420201PNS003STSPLT0001.PDF	EPIC SOURCE TIMESERIES PLOT	
P0303420201PNS003SRCREG0001.ASC	EPIC SOURCE DS9 REGION	
P0303420201PNS003SRCARF0001.FTZ	EPIC ANCILLARY RESPONSE FUNCTION	

Source Dec (Hex)	Inst.	Exp. Id
1 (001)	EMOS1	S001
1 (001)	EMOS2	S002
1 (001)	EPN	S003
2 (002)	EMOS1	S001
2 (002)	EMOS2	S002
2 (002)	EPN	S003
3 (003)	EMOS1	S001

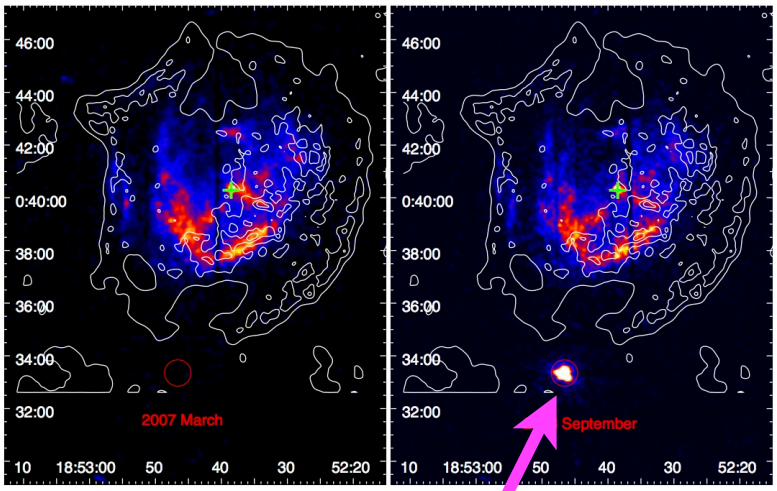
13 30 7.48	47 11 6.4	0.2	667.9 ±29.5	3.13e-13 ±2.58e-14	6822	1367	25012	0	FFTFFFFFFFFF	Y

Importance of archive + S/W analysis capabilities

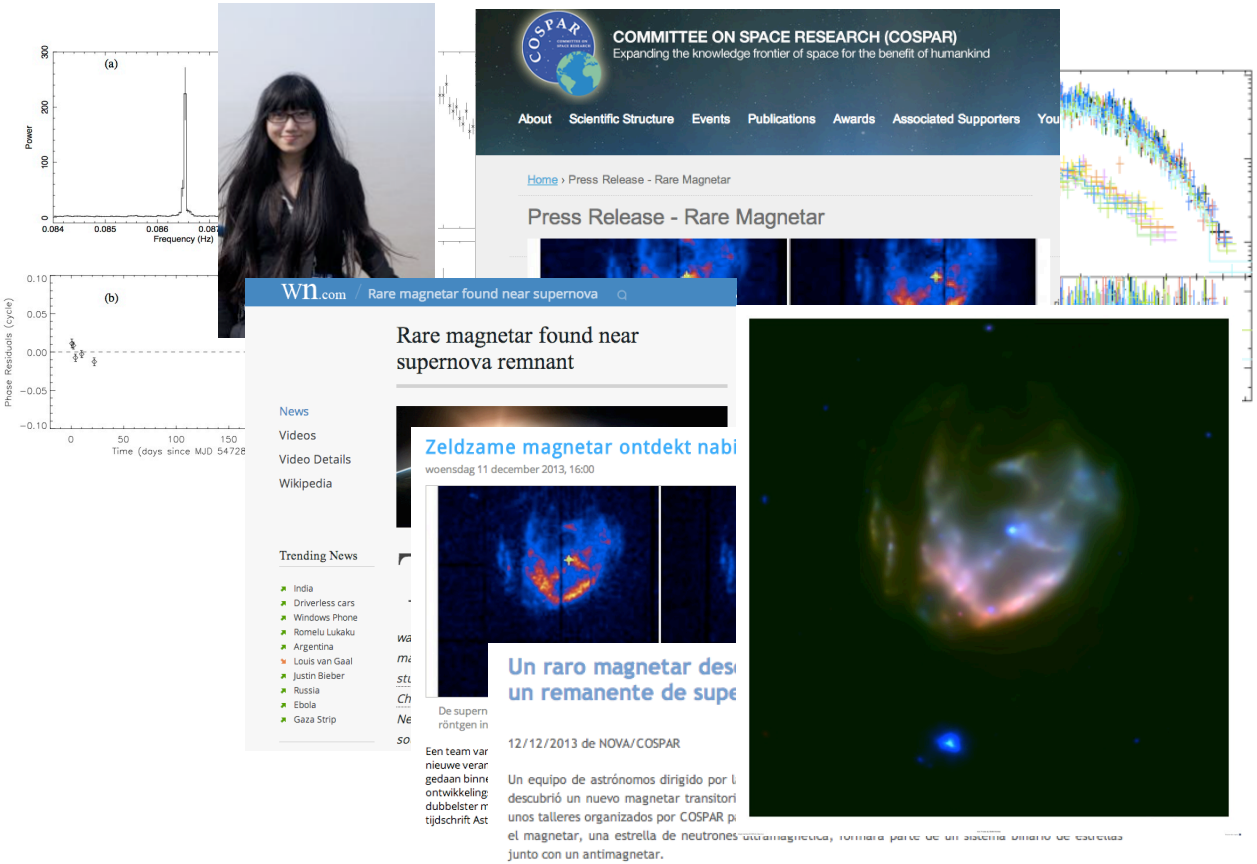
Is **deep** access to data needed after a mission is over?

- data flood - only a certain percentage of the data acquired gets properly analysed
- data analysis tools are a major part of archival research
- data combination ... data extension ... evolution of processing techniques ...
- ...

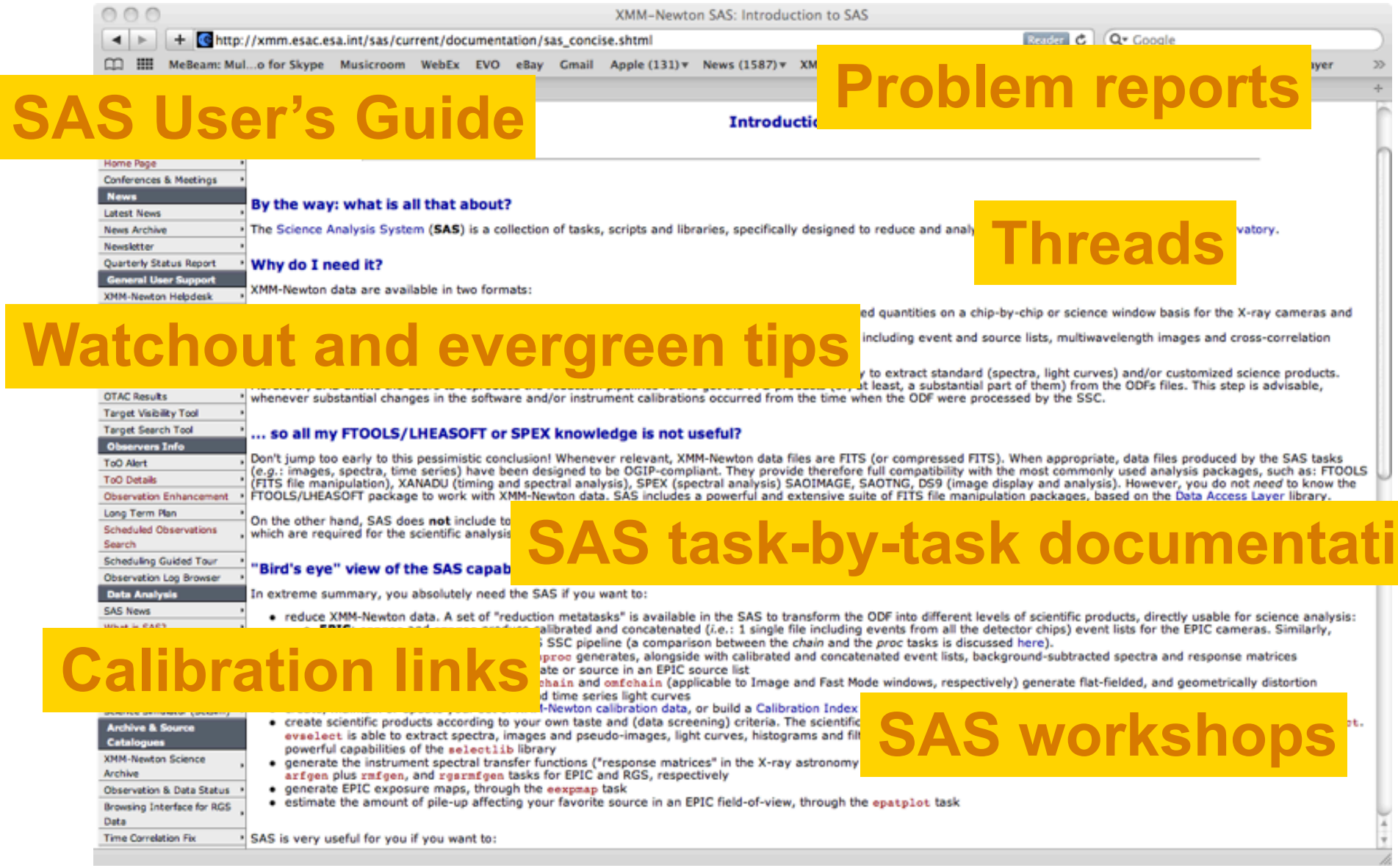
Transient magnetar discovered
in archival data / identified in a COSPAR
Workshop by a student 5 years after data
taken & published



3XMM J185246.6+003317



SAS documentation



What I need to do ...

… before starting to analyse data of an XMM-Newton observation:

1. Verify the quality of the pre-processed scientific products (PPS), produced by the automatic Pipeline processing

2. Check the expected accuracy of the XMM-Newton calibrations, through:
   - Instrument calibration status reports
   - SAS Science Validation Reports
   - Current Calibration File (CCF) Release Notes

3. Compare your own set of calibration files with the latest available
   - Reduce the data again if a calibration file has changed, which may affect your scientific conclusions. Always stay on the safe side!

4. Once you have installed SAS, your job is not finished ...
   - Check the SAS "watchout and evergreen" SAS pages, which contain known caveats or bugs
   - Subscribe to the calibration mailing list
   - Install an automatic mirror of the calibration files
   - Make use of the threads, would you like to learn something new

Shaping the SAS future: RISA

So far, SAS runs locally on user's machine:

Integration on several different platforms + distribution
Large maintenance due to need of compatibility with new libraries in new versions — SOC
SAS download + installation + setup necessary
Data + Calibration DB download — User

Running SAS through Web Services?

reduction of maintenance due to the limitation to few platforms (1 ?)
easy to be fully "frozen" from a certain point in time
neither SAS installation nor data download needs by single user
automatic access to large H/W and S/W resources (ESAC Grid + VO tools)
full data access (processing close to XSA and central CCF repository)

Further Advantages:

- processing in semi-batch mode large amounts of data
- data combination
- complement for archive >> on-the-fly reprocessing
- size-able according to needs - scalability

Disadvantage:

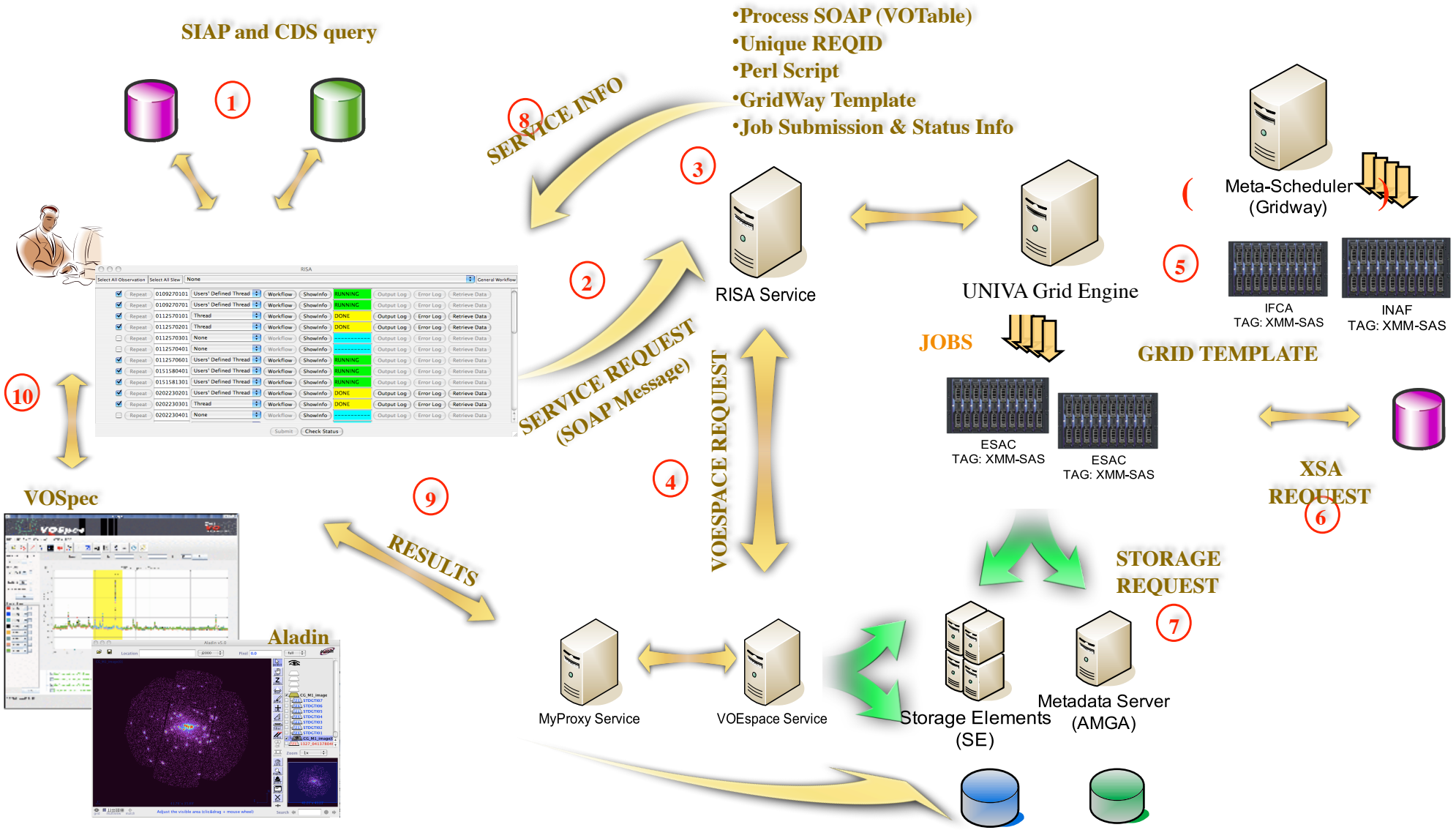
- no own scripting possible ...
  "only" SAS workflows

but

... FTOOLS could be added
... other "certified" S/W

>> complemented with VM (based on same OS)
+ longer cycles / possibility of freezing for long periods
+ allowing for larger control of analysis, scripting + mixing with other tools (IDL, etc)

Remote Interface for Science Analysis (RISA)



Working with RISA

Define workflow, set SAS tasks parameter

Search data

Let the GRID work (Web Service + GridWay)

XSA Access (AIOClient)

Choose data, get info (SIAP)

Start web service

"VO"Access results, retrieve data for local work (VOEspace)

RISA: Remote Interface for Science Analysis

RISA β version released years ago... but not yet public

- showing full analysis capabilities, full access to all parameters
  massive processing capability << ---- >> highest granularity

- several configurable “pipelines” integrated in the system

- processing start at ODF or PPS level

- documentation written as “threads”

Main lines of development:

- services to be integrated into the next XSA version (on-the-fly reprocessing, +++)
- more user friendly / intuitive GUI
- “light” RISA services (mini-pipelines for fast response)
- combination of observations

Advantages:

- reduction of maintenance to single OS
- ideal for processing large amounts of data
- complement for archive >> on-the-fly reprocessing
- sizeable according to needs

>> complemented with VM (based on same OS)
+ longer cycles / possibility of freezing for long periods
+ allowing for larger control of analysis, scripting + mixing with other tools (IDL, etc)